Beyond Circles

Stationary axisymmetric black holes and the breaking of circularity

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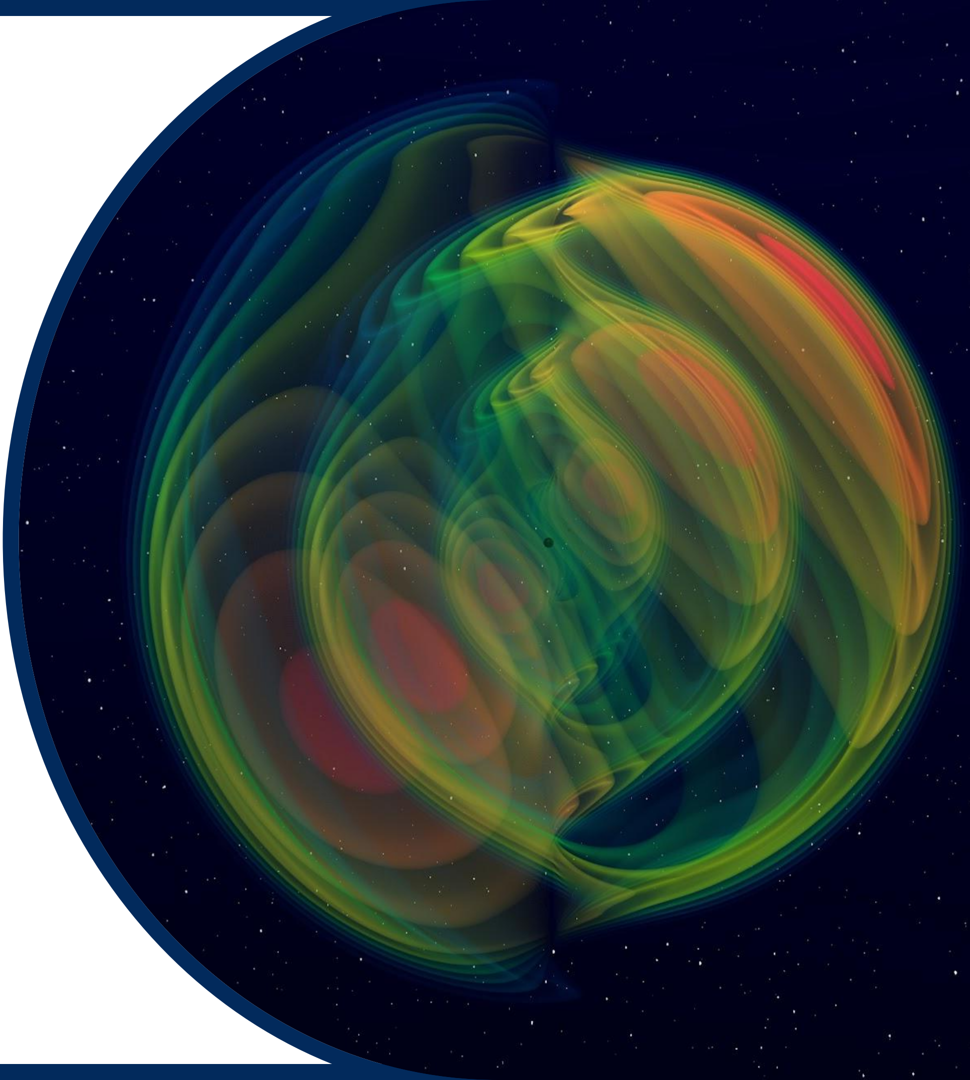


Irène Joliot-Curie
Laboratoire de Physique des 2 Infinis

The talk, in a nutshell

Topic

stationary & axisymmetric isolated objects
[i.e. rotating BHs]

‘Claim’

many well-established properties are not obvious beyond GR

CIRCULARITY

‘additional’ symmetry of BHs in (vacuum) GR

- what does it imply?
- what do we lose by not having it?

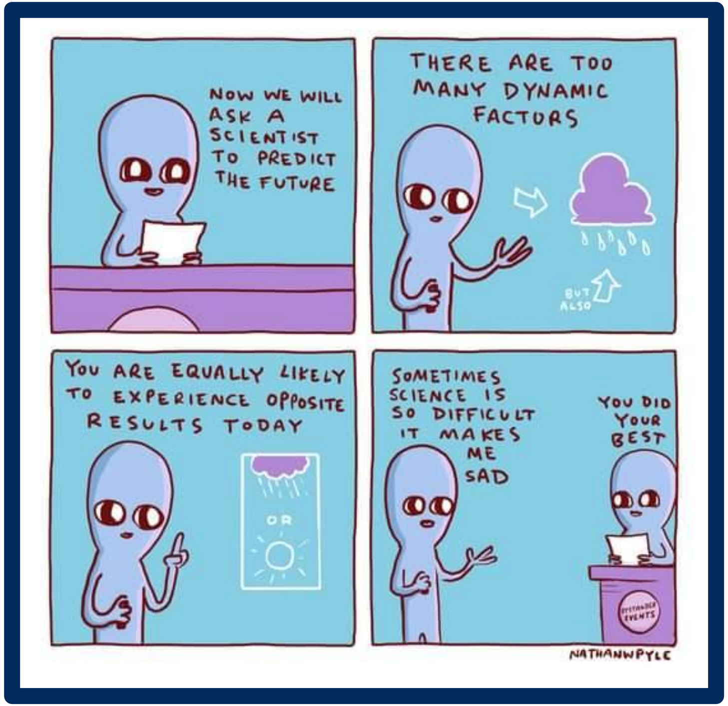
[Keenan Crane [link]]

Outline

- What is circularity?
  - Definition(s) and interpretation
  - Why circularity?
  - Why *not* circularity?
- Moving beyond circularity
  - A metric of sufficient generality
    [*i.e. the most general rotating metric*]
  - 'Solving' circularity
- Properties of non-circular BHs
  - Rotosurface
  - Horizon mechanics
  - Towards thermodynamics

E. Babichev, **JM,** JCAP 10 (2025) 011 [2505.08880]

F. Del Porro, **JM,** [2511.02911]



[Nathan W. Pyle]

What is circularity?

Def

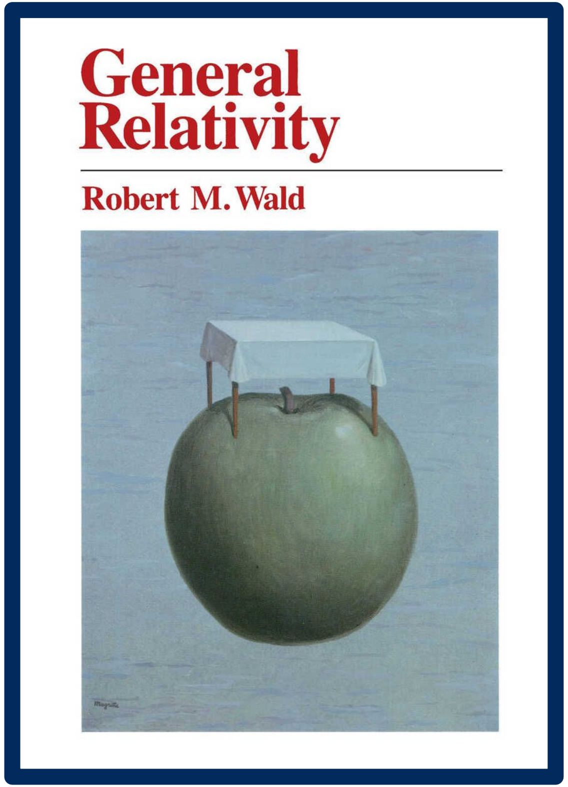
A stationary & axisymmetric spacetime
is said to be *circular*
if it can be foliated by codimension-2
surfaces orthogonal to the Killing vectors

Killing vectors
ξ^μ ……… stationarity
ψ^μ ……… axisymmetry

by Frobenius' theorem

$$\xi_{[\mu}\psi_\nu\partial_\rho\xi_{\sigma]} = 0$$
$$\xi_{[\mu}\psi_\nu\partial_\rho\psi_{\sigma]} = 0$$

[everywhere]



Circularity and geometry

Circularity conditions

$$\xi_{[\mu}\psi_{\nu}\partial_{\rho}\xi_{\sigma]} = 0$$
$$\xi_{[\mu}\psi_{\nu}\partial_{\rho}\psi_{\sigma]} = 0$$

geometric statement on Killing flows

Practically,
there exist coordinates such that

$$g_{\mu\nu} = \begin{pmatrix} * & 0 & 0 & * \\ 0 & * & 0 & 0 \\ 0 & 0 & * & 0 \\ * & 0 & 0 & * \end{pmatrix} \begin{matrix} t \\ r \\ \theta \\ \varphi \end{matrix}$$

t, φ.... Killing coordinates

$(t-\varphi)$ symmetry:
$t \to -t$, $\varphi \to -\varphi$

5 'free' functions,
$g_{\mu\nu} = g_{\mu\nu}(r,\theta)$

i.e. Boyer–Lindquist coordinates

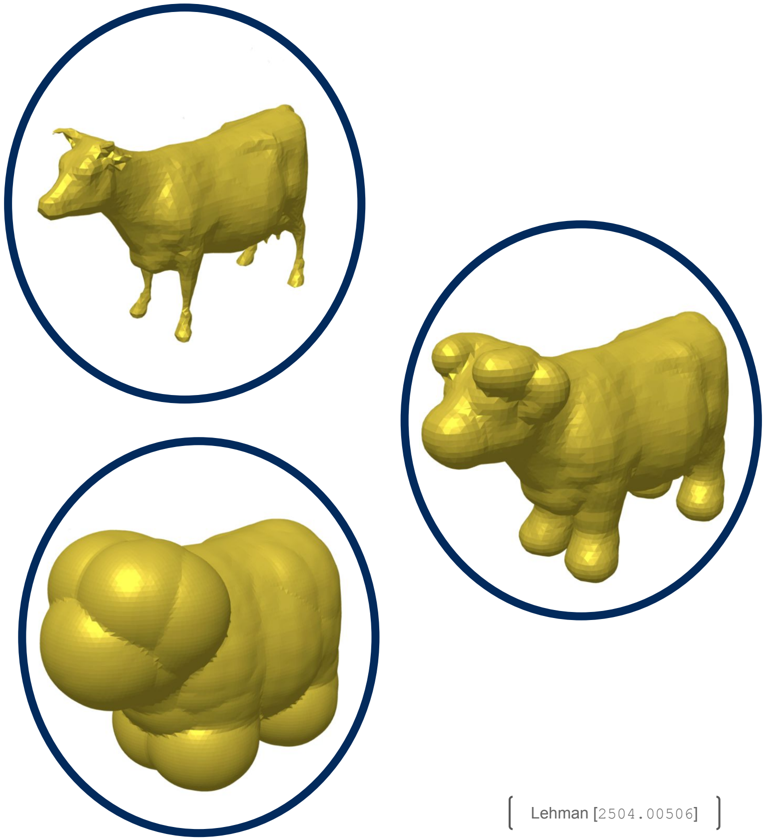
Why circularity?

Circularity is quite ubiquitous, because:

- many GR sol's are circular
- some beyond-GR sol's are circular
- *'square peg in a circular hole'*: beyond-GR sol's that are continuously connected to GR are circular

  [Xie, Zhang, Silva, de Rham, Witek, Yunes, PRL 126 (2021) 241104 [2103.03925]]

- simplicity
  - block-diagonalisation of $g_{\mu\nu}$
  - and more…



[Lehman [2504.00506]]

Why *not* circularity?

Some beyond GR sol's are *not* circular

- Cubic Galileon [numerical] — Grandclement, CQG 41, 025012 (2024) [2308.11245]; Van Aelst, Gourgoulhon, Grandclément, Charmousis, CQG 37 035007 (2020) [1910.08451]
- Semiclassical gravity [numerical] — Fernandes, PRD 108 (2023) 6, 6 [2305.10382]
- Einstein–æther [numerical] — Adam, Figueras, Jacobson, Wiseman, Class. Quant. Grav. 39 (2022) 12, 125001 [2108.00005]
- Disformal Kerr [analytical] — Anson, Babichev, Charmousis, Hassaine, JHEP 2021 (2021) 18 [2006.06461]; Ben Achour, Liu, Motohashi, Mukohyama, Noui, JCAP 11 (2020) 001 [2006.07245]
- 'Locality principle' regular BHs [analytical-ish] — Eichhorn, Held, JCAP 05 (2021) 073 [2103.13163]

Moving beyond circularity

To explore consequences of circularity breaking,

need ‘new’ *framework*

1. **‘need a metric’:**

   need gauge choice that combines generality and minimality
   [i.e. equivalent of Boyer–Lindquist gauge]

   [E. Babichev, **JM**, JCAP 10 (2025) 011 [2505.08880]]

2. then it depends, we’ll see…

   e.g. for thermodynamics, maybe need new ways of computing surface gravity, etc.

   [F. Del Porro, **JM**, [2511.02911]]

General stationary & axisymmetric metric

What is *the most general* stationary and axisymmetric metric?

No one really knows…

[our answer]

Kerr-like gauge

‘Kerr’ as in ‘Kerr ingoing coordinates’
[not Kerr solution]

$$g_{\mu\nu} = \begin{pmatrix} \circledast & \circledast & 0 & \circledast \\ * & 0 & 0 & \circledast \\ 0 & 0 & \circledast & 0 \\ * & * & 0 & \circledast \end{pmatrix}$$

$v \quad r \quad \theta \quad \varphi$

Killing coordinates

non-Killing coordinates

6 ‘free’ functions

$$\begin{cases} g_{vv}(r,\theta) \\ g_{vr}(r,\theta) \\ g_{r\phi}(r,\theta) \\ g_{\phi\phi}(r,\theta) \\ g_{\theta\theta}(r,\theta) \end{cases}$$

existence not obvious, [see paper]

gauge fixing not complete

Comments:

1. Important open problem!
2. Beware of claims saying that BL-like metrics are most general: no!
3. Current phenomenological parametrisations all circular by desing

‘Solving’ circularity

Kerr-like gauge describes
circular and *non-circular* spacetimes

Idea: use it to ‘solve’ circularity conditions

$$\xi_{[\mu}\psi_\nu\partial_\rho\xi_{\sigma]} = 0$$
$$\xi_{[\mu}\psi_\nu\partial_\rho\psi_{\sigma]} = 0$$

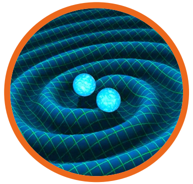
↔

$$g^{vr} = f(r)g^{rr}$$
$$g^{r\phi} = h(r)g^{rr}$$

used them to construct two examples,
both simple deformations of Kerr

‘Minimal’ breaking

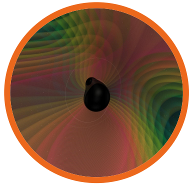
$$g^{vr} = g^{vr}_{\text{Kerr}} + \delta(r,\theta)$$



E. Babichev, **JM**, JCAP 10 (2025) 011 [2505.08880]

‘Not-so-minimal’ breaking

$$M \mapsto m(r,\theta)$$



Properties of non-circular BHs

1. Non-separability of geodesics (and Klein–Gordon), no Killing tensor

not really *due to* circularity breaking

2. (Near-)horizon structure is different

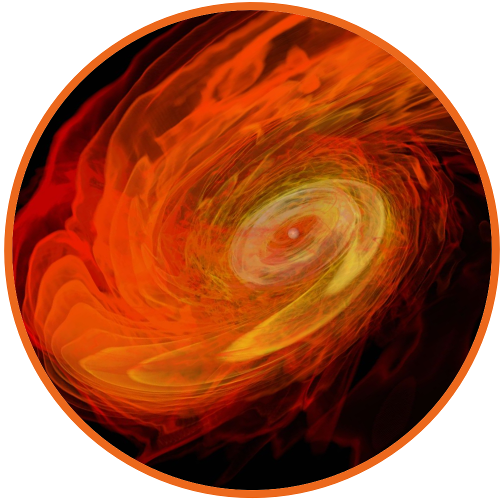
new relevant surface outside *H*: rotosurface

phenomenological consequences
?

event horizon not Killing horizon

thermodynamics
?

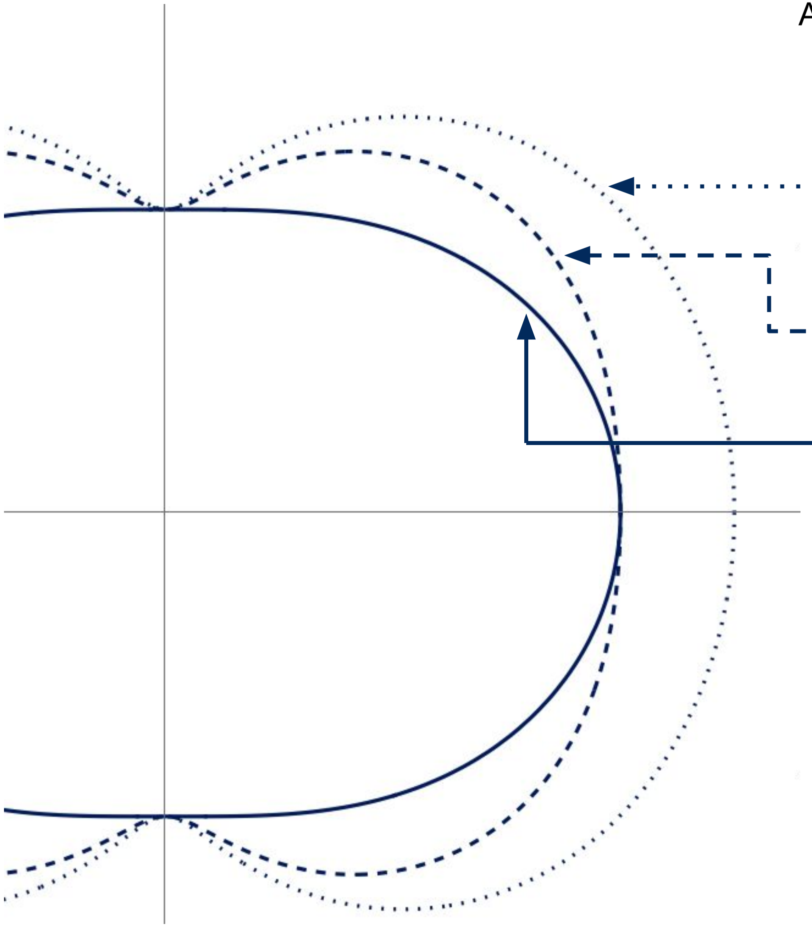
F. Del Porro, **JM,** [2511.02911]



[NASA/AEI/ZIB/M. Koppitz and L. Rezzolla]

The rotosurface

A consequence of circularity breaking: a 'new' surface appears



$$\xi_\mu \xi^\mu = 0$$

ergosurface,
static limit

$$(\xi_\mu + \Omega\, \psi_\mu)(\xi^\mu + \Omega\, \psi^\mu) = 0$$

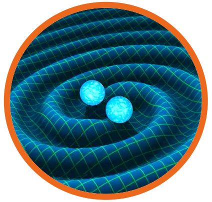
rotosurface,
stationary limit

event
horizon

in Kerr they coincide but in
general they are different

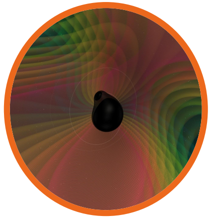
pheno consequences
of rotosorfaces?

Rotosurface and examples



Minimal

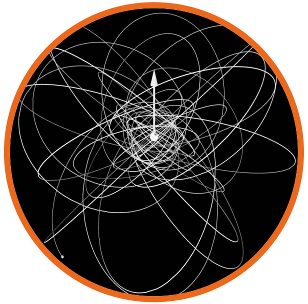
Rotosurface and horizon coincide,
horizon is a ‘sphere’



Non-minimal

Rotosurface and horizon differ,
Everything known analytically

Have fun!



[Maarten van de Meent [link]]

The horizon

Usually
[i.e. vacuum GR]
BH event horizons are also **Killing horizons**

[rigidity theorem]

i.e. they are generated by Killing vectors

Several consequnces:

- Ω_H constant over H
- surface gravity κ constant over H
- different notions of surface gravity coincide

Four laws of 'BH mechanics' ↔ Four laws of 'BH thermodynamics'

0th law $\kappa = \text{const.}$

1st law $\delta M = \frac{\kappa}{8\pi G}\delta A_H + \Omega_H \delta J$

2nd law $\delta A_H \geq 0$

3rd law $\kappa = 0$ unattainable

$\kappa \leftrightarrow T$
$A_H \leftrightarrow S$
$M \leftrightarrow E$

[**with circularity**]

Non-Killing horizon

without circularity

Horizons not Killing, in general

~~Usually~~
~~[i.e. vacuum GR]~~
~~BH event horizons are also **Killing horizons**~~

~~rigidity theorem~~

~~i.e. they are generated by Killing vectors~~

Several consequnces:

- ~~Ω_H constant over H~~
- ~~different notions of surface gravity coincide~~ ?
- ~~surface gravity κ constant over H~~

Four laws of 'BH mechanics' ? ⟷ Four laws of 'BH thermodynamics' ??

0th law $\kappa = \text{const.}$

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3rd law $\kappa = 0$ unattainable

$\kappa \leftrightarrow T$???
$A_H \leftrightarrow S$
$M \leftrightarrow E$

F. Del Porro, **JM**, [2511.02911]

Surface gravities

We computed three notions of surface gravity

F. Del Porro, **JM**, [2511.02911]

κ inaffinity

$$\zeta^\nu \nabla_\nu \zeta^\mu \Big|_{r=H} = \kappa_{\rm i} \zeta^\mu \Big|_{r=H}$$

κ normal

$$\nabla_\mu \left(\zeta_\nu \, \zeta^\nu\right) \Big|_{r=H} = 2\kappa_{\rm n} \zeta_\mu \Big|_{r=H}$$

κ peeling

[see next slide]

on-horizon surface gravities

near-horizon surface gravity

$$\kappa_{\rm i}(\theta) \neq \kappa_{\rm n}(\theta) = \kappa_{\rm p}(\theta)$$

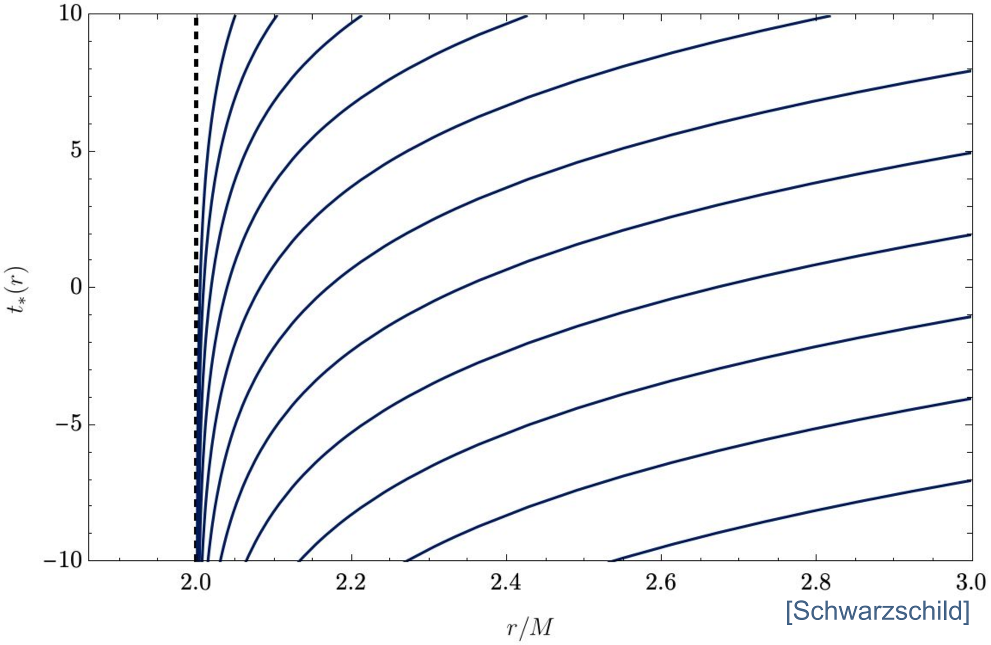
non-constant, and different

assume the horizon is $r = H(\theta)$

then define $\zeta_\mu := \frac{\alpha}{g^{vr}} \partial_\mu \left[r - H(\theta)\right]$

Peeling surface gravity

Outgoing causal geodesics "peel off" of the horizon



[Schwarzschild]

$$t_* \sim \frac{1}{\kappa_{\rm p}} \log(r)$$

peeling surface gravity

despite the lack of geodesic separability, we can compute:

$$\kappa_{\rm p} = \kappa_{\rm n}$$

w/ good choice of normalisation

Hawking radiation

The peeling entails Hawking radiation

Via tunnelling method, we find:
an observer at infinity detects particles with spectral density

$$\langle \hat{\mathcal{N}}_{\omega m} \rangle = \frac{\upsilon(\omega, m)}{\exp\left[\frac{2\pi}{\kappa_{\mathrm{p}}}(\omega - m\sigma_H)\right] - 1}$$

Bose–Einstein-ish
distribution w/

$$T_H = \frac{\kappa_{\mathrm{p}}}{2\pi} \quad \text{and} \quad \mu_H = m\sigma_H$$

AI generated

Towards thermodynamics

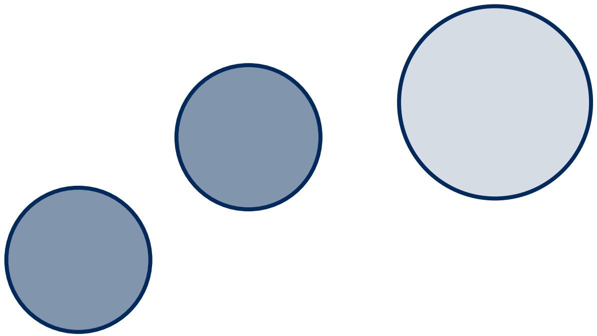
Comments on the four laws of mechanics of non-Killing horizons:

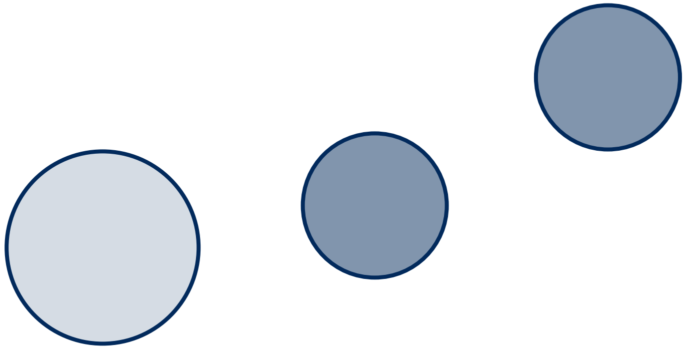
[0th law] constancy of κ: does not hold

[1st law] energy conservation: [w/ caveats] holds in averaged form

[2nd law] horizon's area cannot decrease: holds under the same assumptions as usual

[3rd law] unattaibility of extremal state: same as usual





Open questions:

- Thermodynamic interpretation?
- Out of equilibrium / local (i.e. non global) equilibrium / ?
- [...]

Recap & Food for thought

Circularity
'accidental' symmetry

[might not hold beyond GR]

1. found most general stationary & axisymmetric metric
2. solved circularity conditions
3. constructed simple examples

E. Babichev, **JM,** JCAP 10 (2025) 011 [2505.08880]

main consequences

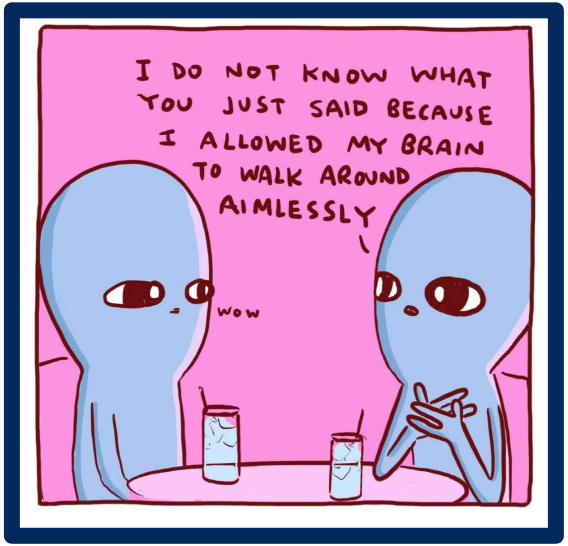
rotosurface separate from horizon

event horizon not Killing

surface gravities are different and not constant

yet, Hawking radiation; thermodynamics?

F. Del Porro, **JM,** [2511.02911]



[Nathan W. Pyle]

What's next?

What about slow rotation?

Usually, slow rotation means Hartle–Thorne [Lense–Thirring] metric

$$ds^2 = -f(r)\mathrm{d}t^2 + \tilde{f}(r)\mathrm{d}r^2 + r^2\left[\mathrm{d}\theta^2 + \sin^2(\theta)\mathrm{d}\phi^2\right] - 2r^2\sin^2(\theta)\Omega(r,\theta)\mathrm{d}t\mathrm{d}\phi$$

this is circular

[$\Omega(r,\vartheta) = \mathcal{O}(a)$]

$$ds^2 = -f(r)\mathrm{d}v^2 + 2h(r)\mathrm{d}v\mathrm{d}r + r^2\left[\mathrm{d}\theta^2 + \sin^2(\theta)\mathrm{d}\phi^2\right] - 2r^2\sin^2(\theta)\left[\Omega(r,\theta)\mathrm{d}t\mathrm{d}\phi + \Lambda(r,\theta)\mathrm{d}r\mathrm{d}\phi\right]$$

e.g. this is not circular

[$\Omega(r,\vartheta)$, $\Lambda(r,\vartheta) = \mathcal{O}(a)$]

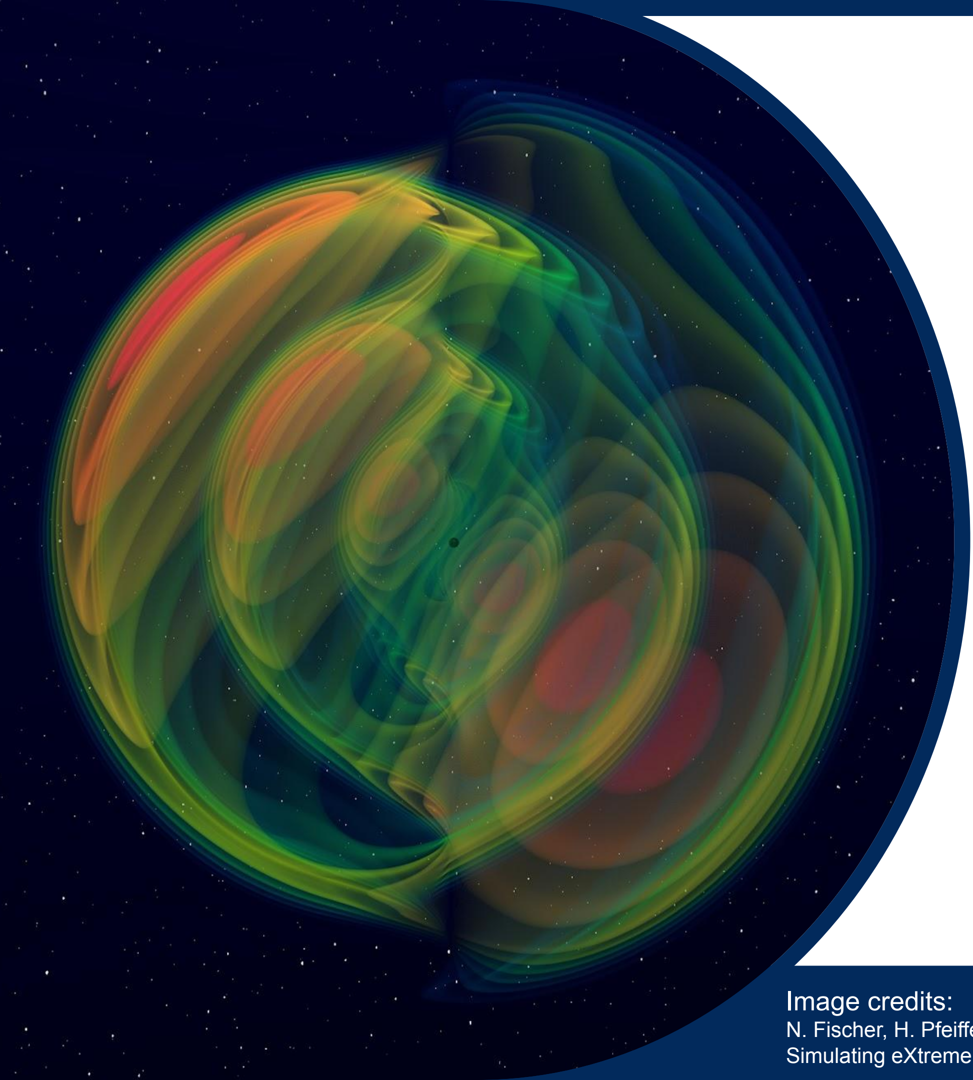
soon on (non-)circularity TV

Is there something 'protecting' circularity at $\mathcal{O}(a)$,

or are we missing some slowly rotating sol's?

?

stay tuned…



Thanks!

Get in touch

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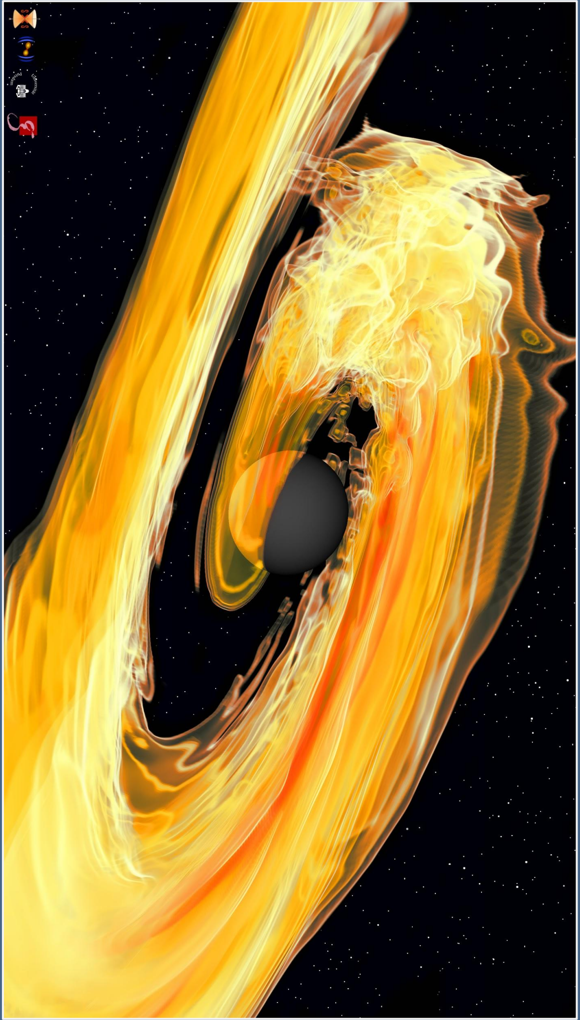
Image credits:
N. Fischer, H. Pfeiffer, A. Buonanno (Max Planck Institute for Gravitational Physics), Simulating eXtreme Spacetimes (SXS) Collaboration

Backup Slides



[BckUp] Circularity and (*more*) geometry

Since you liked…

$$\xi_{[\mu}\psi_\nu\partial_\rho\xi_{\sigma]} = 0$$

$$\xi_{[\mu}\psi_\nu\partial_\rho\psi_{\sigma]} = 0$$

[everywhere]

… you might also like

$$\xi^\mu R_{\mu[\nu}\psi_\rho\xi_{\sigma]} = 0$$

$$\psi^\mu R_{\mu[\nu}\psi_\rho\xi_{\sigma]} = 0$$

[everywhere]

\+

$$\xi_{[\mu}\psi_\nu\partial_\rho\xi_{\sigma]} = 0$$

$$\xi_{[\mu}\psi_\nu\partial_\rho\psi_{\sigma]} = 0$$

[at least at one point]

equivalent formulations
in ‘normal’ situations

(use Killing eq.’s and some assumptions)

[BckUp] Circularity and matter

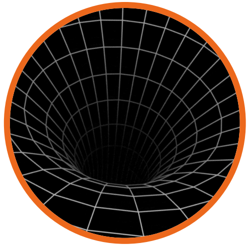
$$\xi^\mu R_{\mu[\nu}\psi_\rho\xi_{\sigma]} = 0$$
$$\psi^\mu R_{\mu[\nu}\psi_\rho\xi_{\sigma]} = 0$$

[everywhere]

[statement on symmetry of sources via field equations]

Einstein spaces automatically circular

[Kerr–Newman (A)dS, etc.]

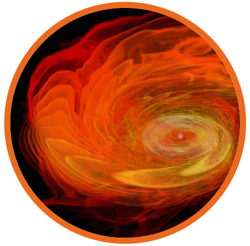


[James Zanoni/Quanta Magazine]

SET of a fluid: velocity

$$u^\mu = \alpha\,\xi^\mu + \beta\,\psi^\nu$$

[$u^{[\mu}\xi^\nu\psi^{\rho]} = 0$]



[NASA/AEI/ZIB/M. Koppitz and L. Rezzolla]

[BckUp] Square peg in circular hole

Y. Xie, J. Zhang, H.O. Silva, C. de Rham, H. Witek and N. Yunes, PRL 126 (2021) 241104 [2103.03925]

'Square peg in a cirucla hole'

beyond-GR theories like

$$\mathcal{L}_0\left(g_{\mu\nu},\varphi\right)+\alpha\mathcal{L}_M\left(g_{\mu\nu},\varphi\right)$$

operators up to dim 4

higher order

that allow perturabtive treatment

$$g_{\mu\nu}=g^{(0)}_{\mu\nu}+\sum_n\alpha^n g^{(n)}_{\mu\nu}$$

$$\varphi=\varphi^{(0)}+\sum_n\alpha^n\varphi^{(n)}$$

theorem

if

φ scalar …… $\xi^\mu\partial_\mu\varphi=0$

φ vector …… φ circular

[…]

then

$g_{\mu\nu}$ circular order by order in α

[BckUp] Kerr-like gauge

Kerr-like gauge condition

$$0 = g_{v\theta} = g_{\tilde{v}\tilde{v}} \frac{\partial V}{\partial \theta} + g_{\tilde{v}\tilde{r}} \frac{\partial R}{\partial \theta} + g_{\tilde{v}\tilde{\theta}} \frac{\partial \Theta}{\partial \theta} + g_{\tilde{v}\tilde{\phi}} \frac{\partial \Phi}{\partial \theta}$$

$$0 = g_{\phi\theta} = g_{\tilde{\phi}\tilde{v}} \frac{\partial V}{\partial \theta} + g_{\tilde{\phi}\tilde{r}} \frac{\partial R}{\partial \theta} + g_{\tilde{\phi}\tilde{\theta}} \frac{\partial \Theta}{\partial \theta} + g_{\tilde{\phi}\tilde{\phi}} \frac{\partial \Phi}{\partial \theta}$$

$$\begin{aligned} 0 = g_{r\theta} = {} & [g_{v\theta}] \frac{\partial V}{\partial r} + [g_{\phi\theta}] \frac{\partial \Phi}{\partial r} \\ & + \left[g_{\tilde{v}\tilde{r}} \frac{\partial R}{\partial r} + g_{\tilde{v}\tilde{\theta}} \frac{\partial \Theta}{\partial r} \right] \frac{\partial V}{\partial \theta} + \left[g_{\tilde{r}\tilde{\phi}} \frac{\partial R}{\partial r} + g_{\tilde{\theta}\tilde{\phi}} \frac{\partial \Theta}{\partial r} \right] \frac{\partial \Phi}{\partial \theta} \\ & + \left[g_{\tilde{r}\tilde{r}} \frac{\partial R}{\partial r} + g_{\tilde{r}\tilde{\theta}} \frac{\partial \Theta}{\partial r} \right] \frac{\partial R}{\partial \theta} + \left[g_{\tilde{r}\tilde{\theta}} \frac{\partial R}{\partial r} + g_{\tilde{\theta}\tilde{\theta}} \frac{\partial \Theta}{\partial r} \right] \frac{\partial \Theta}{\partial \theta} \end{aligned}$$

$$\begin{aligned} 0 = g_{rr} = {} & g_{\tilde{v}\tilde{v}} \left(\frac{\partial V}{\partial r} \right)^2 + g_{\tilde{r}\tilde{r}} \left(\frac{\partial R}{\partial r} \right)^2 + g_{\tilde{\theta}\tilde{\theta}} \left(\frac{\partial \Theta}{\partial r} \right)^2 + g_{\tilde{\phi}\tilde{\phi}} \left(\frac{\partial \Phi}{\partial r} \right)^2 \\ & + 2 g_{\tilde{v}\tilde{r}} \frac{\partial V}{\partial r} \frac{\partial R}{\partial r} + 2 g_{\tilde{v}\tilde{\theta}} \frac{\partial V}{\partial r} \frac{\partial \Theta}{\partial r} + 2 g_{\tilde{v}\tilde{\phi}} \frac{\partial V}{\partial r} \frac{\partial \Phi}{\partial r} \\ & + 2 g_{\tilde{r}\tilde{\theta}} \frac{\partial R}{\partial r} \frac{\partial \Theta}{\partial r} + 2 g_{\tilde{r}\tilde{\phi}} \frac{\partial R}{\partial r} \frac{\partial \Phi}{\partial r} + 2 g_{\tilde{\theta}\tilde{\phi}} \frac{\partial \Theta}{\partial r} \frac{\partial \Phi}{\partial r} \end{aligned}$$

perform coordinate change

[tilt Cauchy surface]

$$r \to r\,, \quad \theta \to \theta + r$$

$$\frac{\partial U_i}{\partial \theta} \to \frac{\partial U_i}{\partial \theta}\,, \qquad \frac{\partial U_i}{\partial r} \to \frac{\partial U_i}{\partial r} + \frac{\partial U_i}{\partial \theta}$$

Cauchy–Kovalvskaya form

$$\frac{\partial U_i}{\partial \theta} = F\left(g_{\tilde{\mu}\tilde{\nu}}\,, \frac{\partial U_i}{\partial r}\,, \partial U_j \right)$$

[BckUp] Remarks

Indeces up	Indeces down
$\frac{g^{vr}}{g^{rr}} = f(r)$	$\frac{g_{vr}g_{\phi\phi} - g_{r\phi}g_{v\phi}}{g_{v\phi}^2 - g_{vv}g_{\phi\phi}} = f(r)$
$\frac{g^{r\phi}}{g^{rr}} = h(r)$	$\frac{g_{r\phi}g_{vv} - g_{vr}g_{v\phi}}{g_{v\phi}^2 - g_{vv}g_{\phi\phi}} = h(r)$
nice	not (so) nice

Solved conditions are existence conditions for coordinate change to ‘Boyer–Lindquist form’

$$g_{\mu\nu} = \begin{pmatrix} g_{vv} & 0 & 0 & g_{v\phi} \\ 0 & g_{rr} & 0 & 0 \\ 0 & 0 & g_{\theta\theta} & 0 \\ * & 0 & 0 & g_{\phi\phi} \end{pmatrix}$$


[NASA, ESA, CSA, STScI (JWST)]

[BckUp] 'Minimal' deformation

Solved circularity conditions

$$g^{vr} = f(r) g^{rr}$$
$$g^{r\phi} = h(r) g^{rr}$$

'soft' breaking

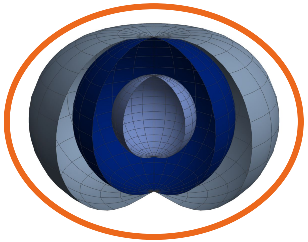
$$g^{vr} = f_{\text{Kerr}} g^{rr}_{\text{Kerr}} + f(r,\theta)_{\text{deform}}$$
$$g^{r\phi} = h_{\text{Kerr}} g^{rr}_{\text{Kerr}}$$

$$g^{\mu\nu} = g^{\mu\nu}_{\text{Kerr}} + \frac{\delta(r,\theta)}{\Sigma} \delta^{\mu}_{v} \delta^{\mu}_{\phi}$$
$$g_{\mu\nu} = \big(\text{Quite-A-Mess}\big)_{\mu\nu}$$

surfaces @ same location

$$r_{\text{erg}} = M + \sqrt{M^2 - a^2 \cos^2\theta}$$

$$r_{\text{H}} = r_{\text{rot}} = M + \sqrt{M^2 - a^2}$$



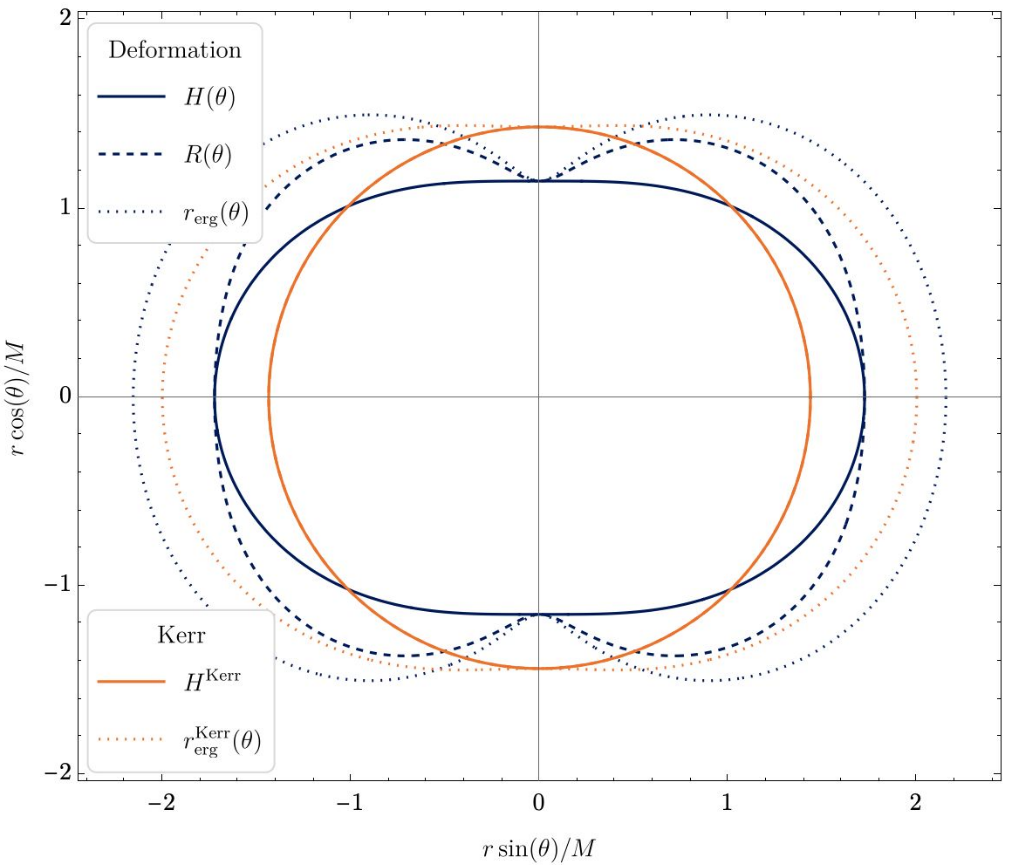
but H not Killing, and

$$\kappa = \left. \frac{r-M}{2rM+\delta(r,\theta)} \right|_{r_H}$$

[BckUp] ‘Not-so-minimal’ deformation

Choose horizon profile *H(θ)*,
then reverse-engineer the metric

horizon, rotosurface, ergosphere
known analytically



[BckUp] Horizon generators

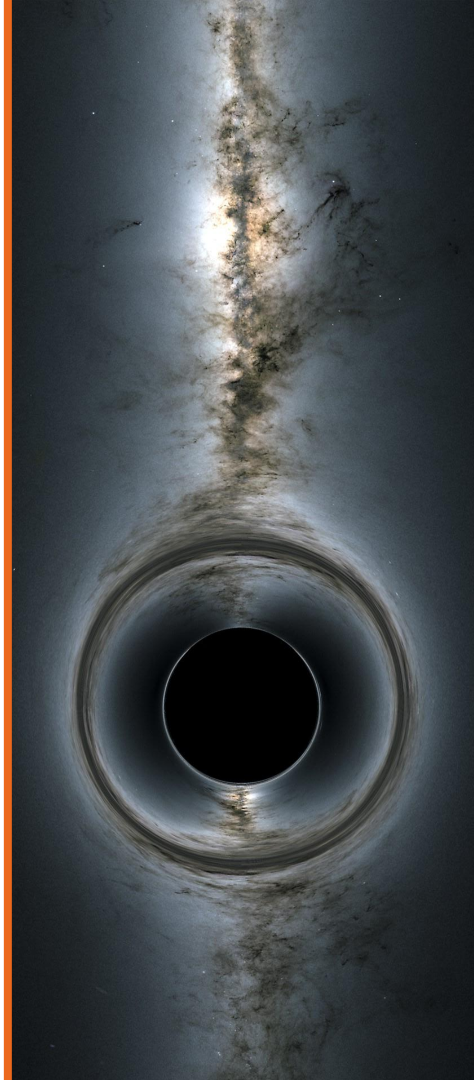
Horizon

- surface: $$r = H(\theta)$$
- normal: $$\zeta_\mu := \frac{\alpha}{g^{vr}} \partial_\mu \left[r - H(\theta) \right]$$
- generator: $$\zeta_\mu \zeta^\mu \Big|_{r=H} = 0$$

$$\zeta^\mu \Big|_{r=H} = \alpha \left[\Xi^\mu - \frac{H' g^{\theta\theta}}{g^{vr}} \tau^\mu \right] \Big|_{r=H}$$

Ξ^μ: would-be Killing vector

τ^μ: tangent to H



[BckUp] Surface gravities

Assume the horizon is $r = H(\theta)$ then define $\zeta_\mu := \frac{\alpha}{g^{vr}} \partial_\mu \left[r - H(\theta) \right]$

inaffinity surface gravity $\zeta^\nu \nabla_\nu \zeta^\mu \Big|_{r=H} = \kappa_\text{i} \zeta^\mu \Big|_{r=H}$

normal surface gravity $\nabla_\mu \left(\zeta_\nu \zeta^\nu \right) \Big|_{r=H} = 2\kappa_\text{n} \zeta_\mu \Big|_{r=H}$

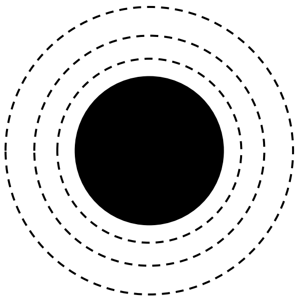
on-horizon surface gravities

compute them both

$$\kappa_\text{n} = \frac{\alpha}{2g^{vr}} \left[\partial_r g^{rr} + (H')^2 \partial_r g^{\theta\theta} \right] \Big|_{r=H}$$

$$\kappa_\text{i} = \kappa_\text{n} + \zeta^\mu \partial_\mu \log(\tfrac{\alpha}{g^{vr}}) \Big|_{r=H}$$

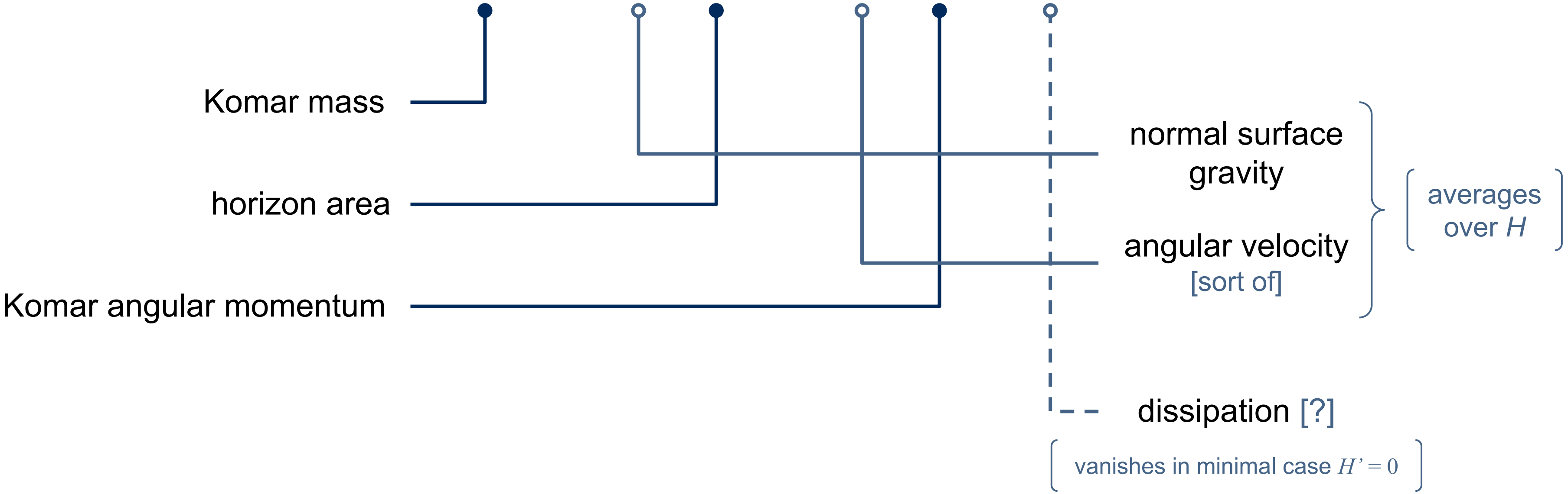
depend on position on *H*

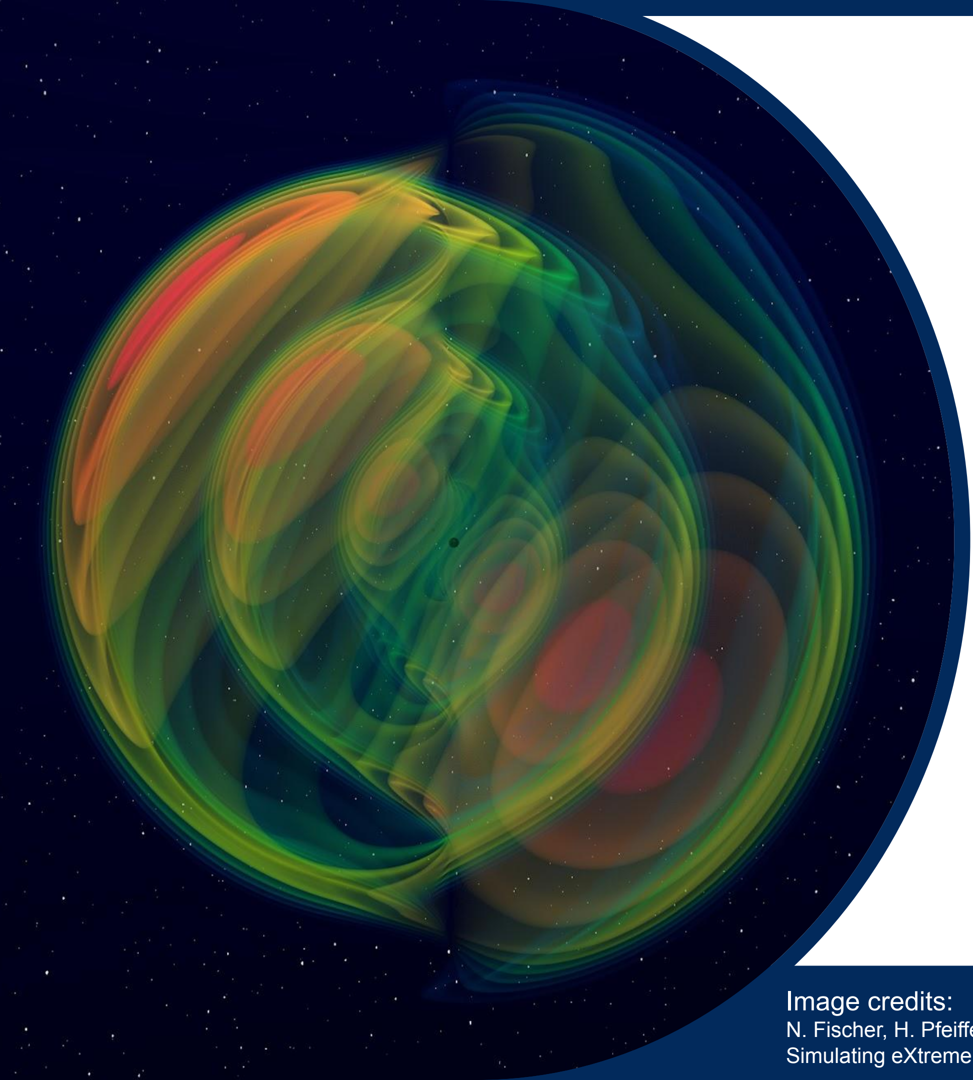


Smarr’s formula

We can prove this geometric identity, no field equations

$$M_H = \frac{\overline{\kappa}}{4\pi G} \mathcal{A}_H + 2\overline{\sigma}_H J_H + \eta$$





Thanks!

Get in touch

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Image credits:
N. Fischer, H. Pfeiffer, A. Buonanno (Max Planck Institute for Gravitational Physics),
Simulating eXtreme Spacetimes (SXS) Collaboration

Friday - 12 Decembre 2025 - https://indico.in2p3.fr/event/37628/overview

Timing: 23 min + 5 min

Beyond circles: stationary axisymmetric black holes and the breaking of circularity

Circularity is an accidental symmetry of the Kerr metric, one that is widely assumed when searching for rotating black hole solutions in modified gravity as well as when constructing models of Kerr mimickers. Though extremely enticing, circularity is often an excessively restrictive assumption, and understanding the consequences of its loss is thus crucially relevant. In this seminar, I wish to present some recent results on the subject: After describing in detail what this symmetry entails, I will show how to construct stationary and axisymmetric spacetimes exhibiting a controlled breaking of circularity; then, I will describe the impact of circularity breaking on the hole's horizon, focusing in particular on the laws of black hole mechanics. This discussion is thus going to be pertinent for anyone with an interest in compact astrophysical objects and their phenomenology, in general relativity and beyond.

The talk, in a nutshell

Many well-established properties of stationary & axisymmetric BHs are not obvious beyond GR

‘form’ of $g_{\mu\nu}$, i.e. coordinate choice

event horizon properties [Killing/not Killing]

CIRCULARITY

Keenan Crane [link]

'Solving' circularity

The circularity conditions are PDEs for metric components

$$\xi_{[\mu}\psi_\nu\partial_\rho\xi_{\sigma]} = 0$$
$$\xi_{[\mu}\psi_\nu\partial_\rho\psi_{\sigma]} = 0$$

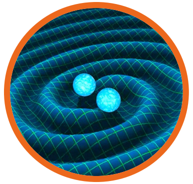
- checking circularity for given metric ---- easy
- constructing 'simple' non-circular metrics ---- hard

idea

'solve' conditions,

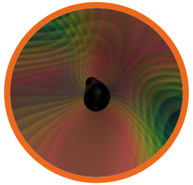
translate them into algebraic relations

$$g^{vr} = f(r) g^{rr}$$
$$g^{r\phi} = h(r) g^{rr}$$



used them to construct two examples, both simple deformations of Kerr:

- 'minimal'
- 'not so minimal'



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