Neutrino emission features from 3D supernova simulations

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Outline

★ Supernova explosion mechanism and hydrodynamical instabilities

★ Detection perspectives adopting first full-scale 3D SN simulations

★ A new instability: Lepton number emission self-sustained asymmetry

★ Conclusions

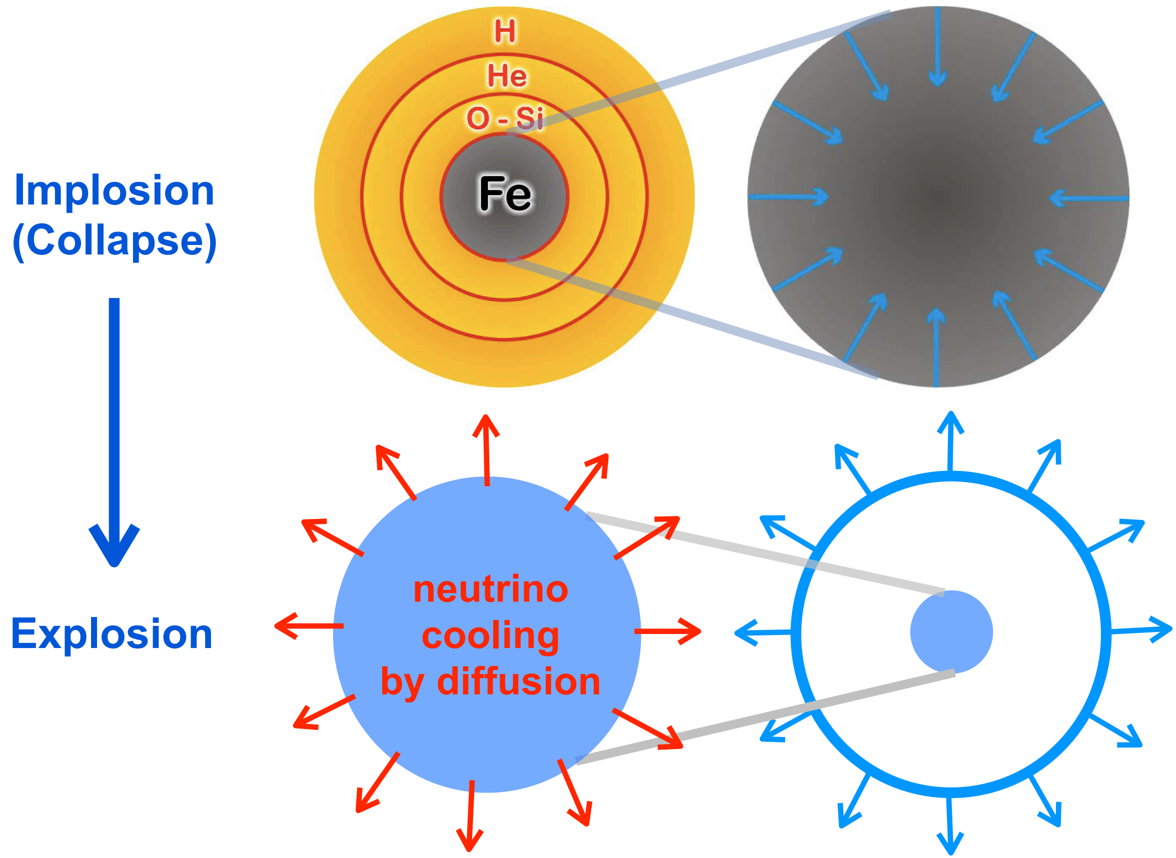
This talk is mainly based on:

- I. Tamborra, F. Hanke, B. Mueller, H.-T. Janka, and G. Raffelt, PRL 111 (2013) 121104.
- I. Tamborra, F. Hanke, H.-T. Janka, B. Mueller, G. Raffelt, and A. Marek, arXiv: 1402.5418.
- I. Tamborra, G. Raffelt, F. Hanke, H.-T. Janka, and B. Mueller, arXiv: 1406.0006.

The Neutrino-driven Explosion Mechanism

Neutrinos in Supernovae

Core-collapse supernovae: Terminal phase of massive stars [$M \geq 8M_{\odot}$].
Stars collapse ejecting the outer mantle by means of shock-wave driven explosions.
Expected rate: 1-3 SN/century in our galaxy (~ 10 kpc).



Neutrinos carry
99% of the released energy
(~ 10^{53} erg).

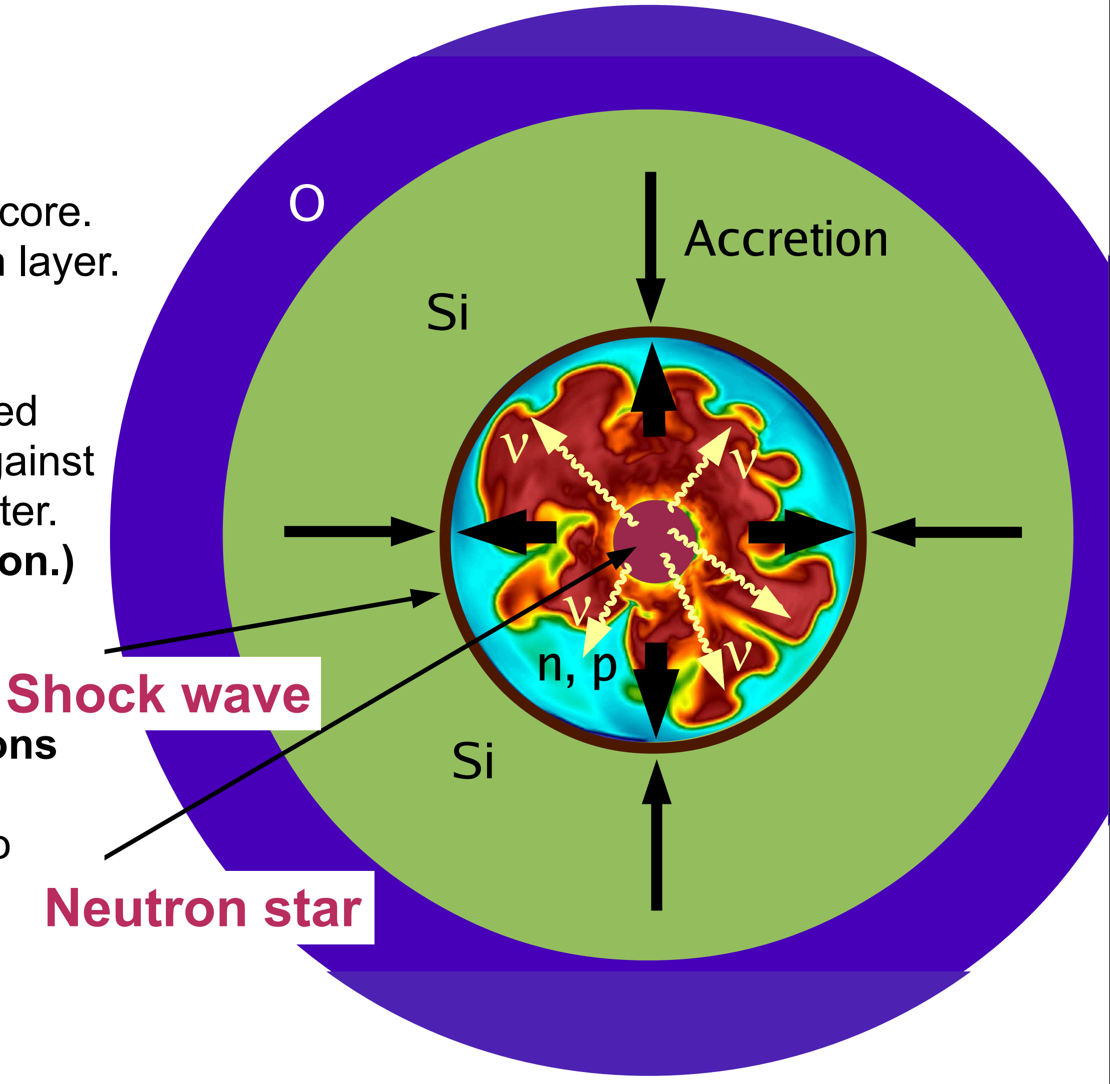
Neutrino typical energies: ~ 15 MeV.
Neutrino emission time: ~ 10 s.

Neutrinos and SN Explosion Mechanism

★ Shock wave forms within the iron core. It dissipates energy dissociating iron layer.

★ **Neutrinos** provide energy to stalled shock wave to start re-expansion against ram pressure of in-falling stellar matter. **(Delayed Neutrino-Driven Explosion.)**

★ **Convection** and **shock oscillations** (standing accretion shock instability, **SASI**) enhance efficiency of neutrino heating and revive the shock.

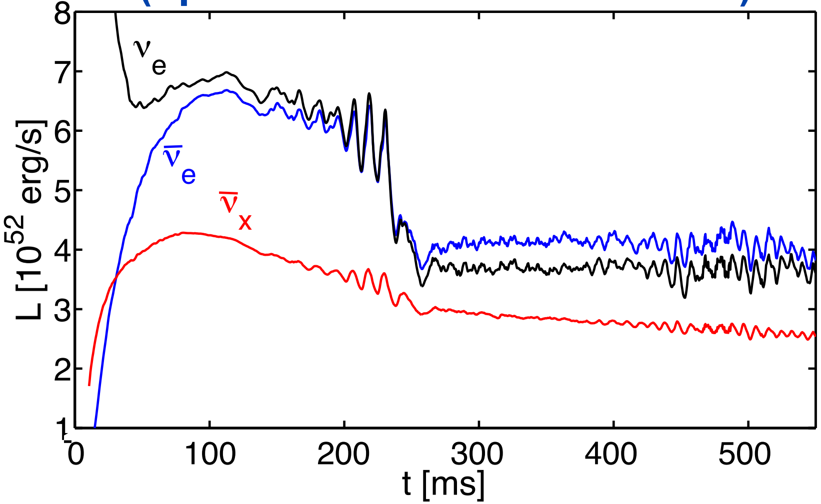


* For more details see H.-T. Janka, arXiv: 1206.2503, H.-T. Janka et al., arXiv: 1211.1378.

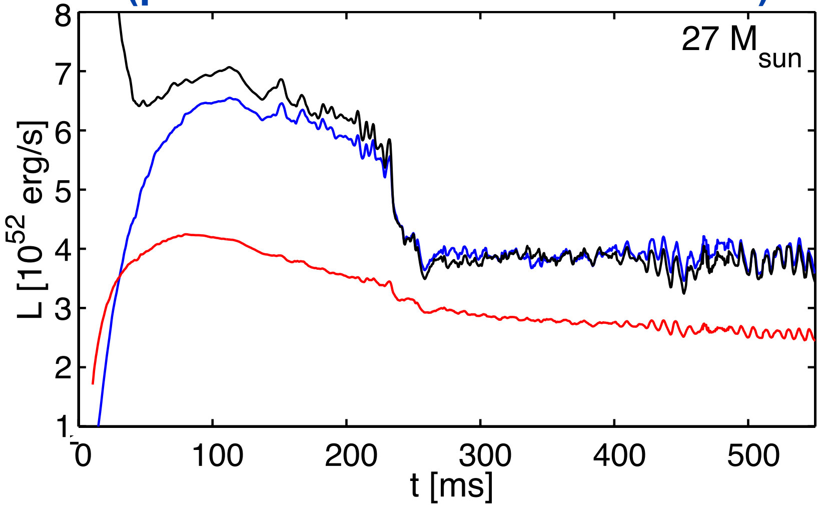
Directional Neutrino Signal

First world-wide 3D SN simulations with detailed neutrino transport available.
SASI and convective motions leave imprints on the neutrino signal.

Close to the SASI plane (optimistic observer direction)



Perpendicularly to the SASI plane (pessimistic observer direction)



Large amplitude modulations close to the plane where spiral SASI mode develops.
Are such modulations detectable?
Are these features generic for any SN progenitor?

For more details see also: F. Hanke et al., ApJ 770 (2013) 66. T. Lund et al., PRD 82 (2012) 063007, PRD 86 (2012) 105031.

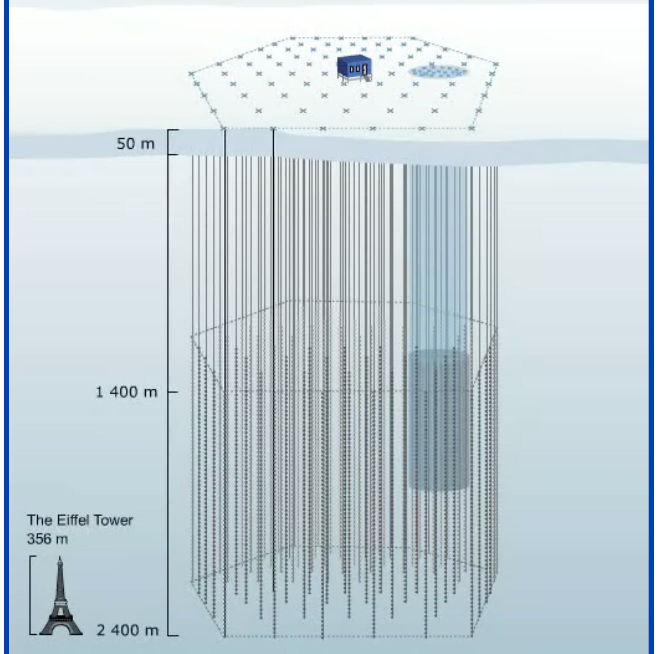
Detection Perspectives

Detection Perspectives

In IceCube and Hyper-Kamiokande, neutrinos are primarily detected by inverse beta decay

$$\bar{\nu}_e + p \rightarrow n + e^+$$

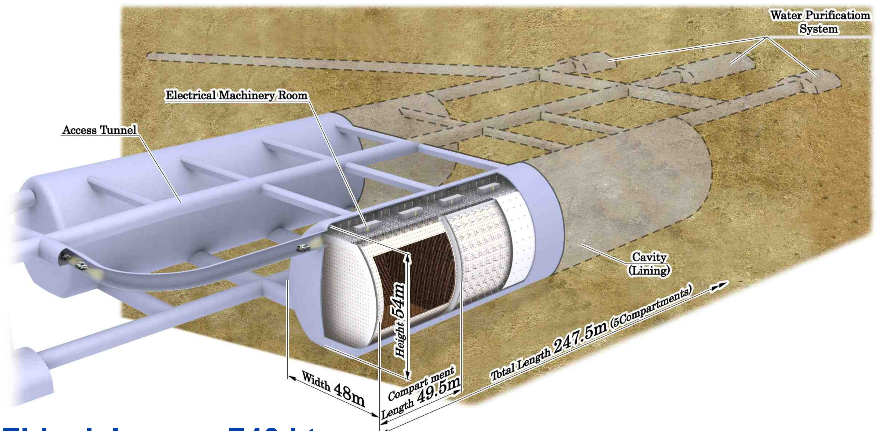
IceCube



Atom

1 km antarctic ice with 5160 PMT
R_{bkgd} = 1480 ms^{-1}

Hyper-Kamiokande



Fiducial mass: 740 kton
Background free signal + event-by-event energy information

* For details see: Abbasi et al., arXiv: 1108.0171 (IceCube), K. Abe et al., arXiv: 1109.3262 (Hyper-K).

SASI Detection Perspectives (27 M_{sun})

Our 3D 27 M_{sun} SN progenitor shows pronounced SASI.
SASI sinusoidal modulation of the neutrino signal will be detectable by IceCube and Hyper-K.

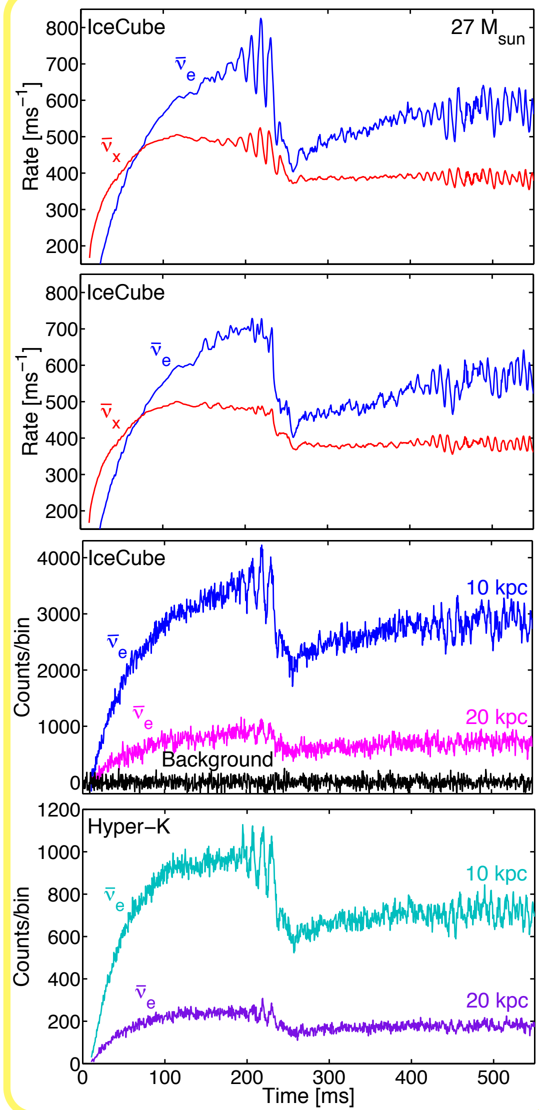
Strong signal modulation (optimistic observer direction)

Weak signal modulation (pessimistic observer direction)

Expected rate above IceCube background

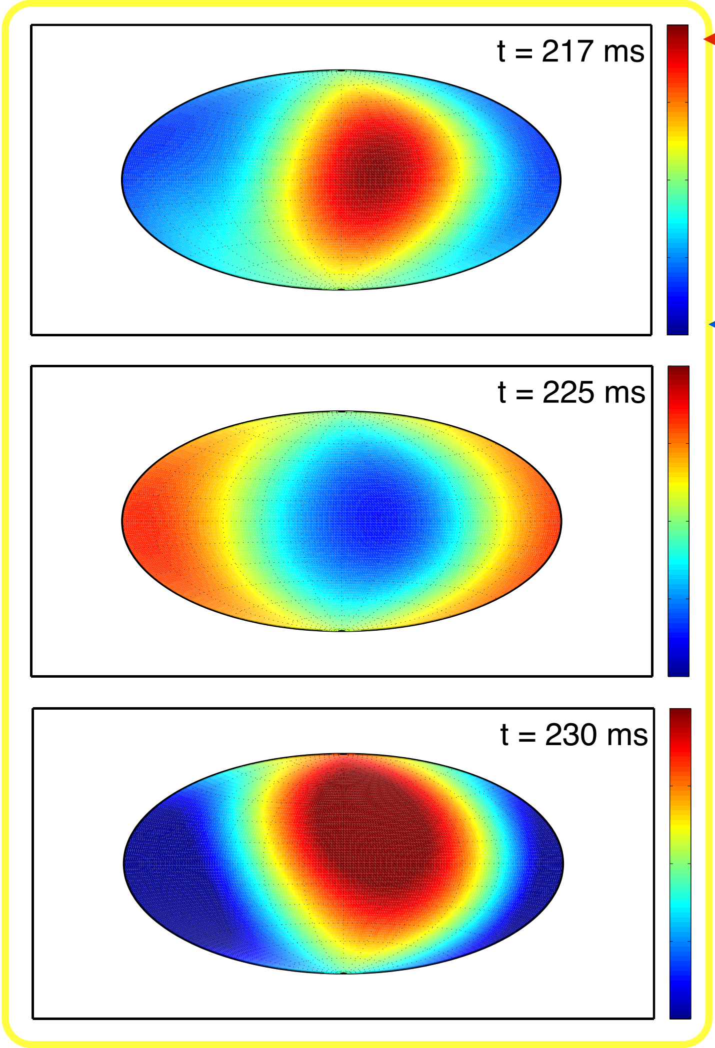
Hyper-K rate = 1/3 IceCube rate

SASI still detectable



SASI Detection Perspectives (27 M_{sun})

Time evolution of the IceCube detection rate on a sky-plot of observer directions.



high detection rate

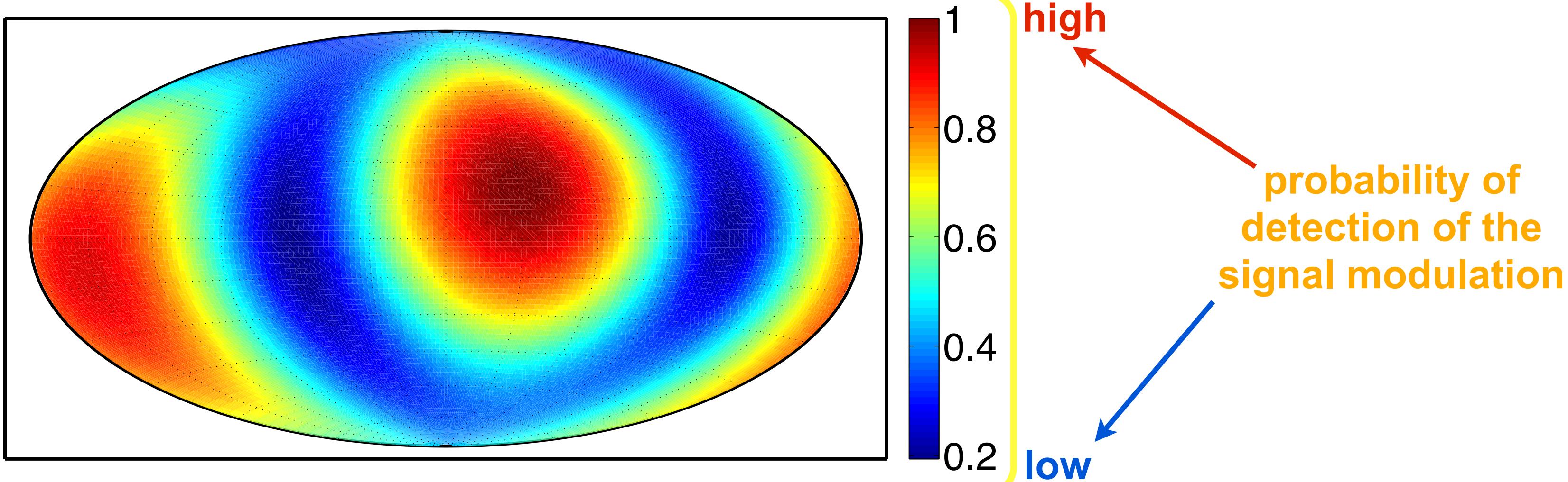
low detection rate

Animated visualization available at:
http://www.mpa-garching.mpg.de/ccsnarchive/data/Hanke2013_movie/index.html

SASI Detection Perspectives (27 M_{sun})

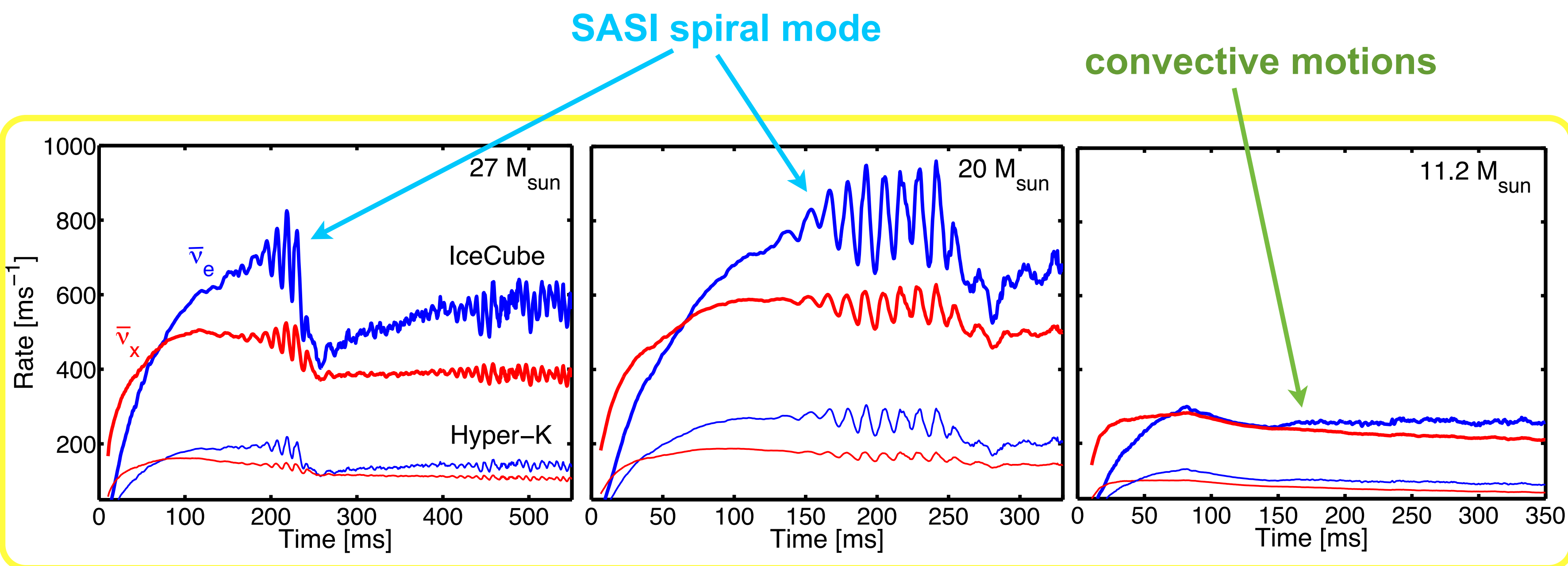
$$\sigma \equiv \left(\int_{t_1}^{t_2} dt \left[\frac{R - \langle R \rangle}{\langle R \rangle} \right]^2 \right)^{1/2}$$

$$[t_1, t_2] = [120, 250] \text{ ms}$$



On average, the fraction of sky where good observation chances apply is significant (> 50%).

Hydro-instability Detection Perspectives



27 M_{sun} SN progenitor:
Two SASI episodes with
convective phase in between.

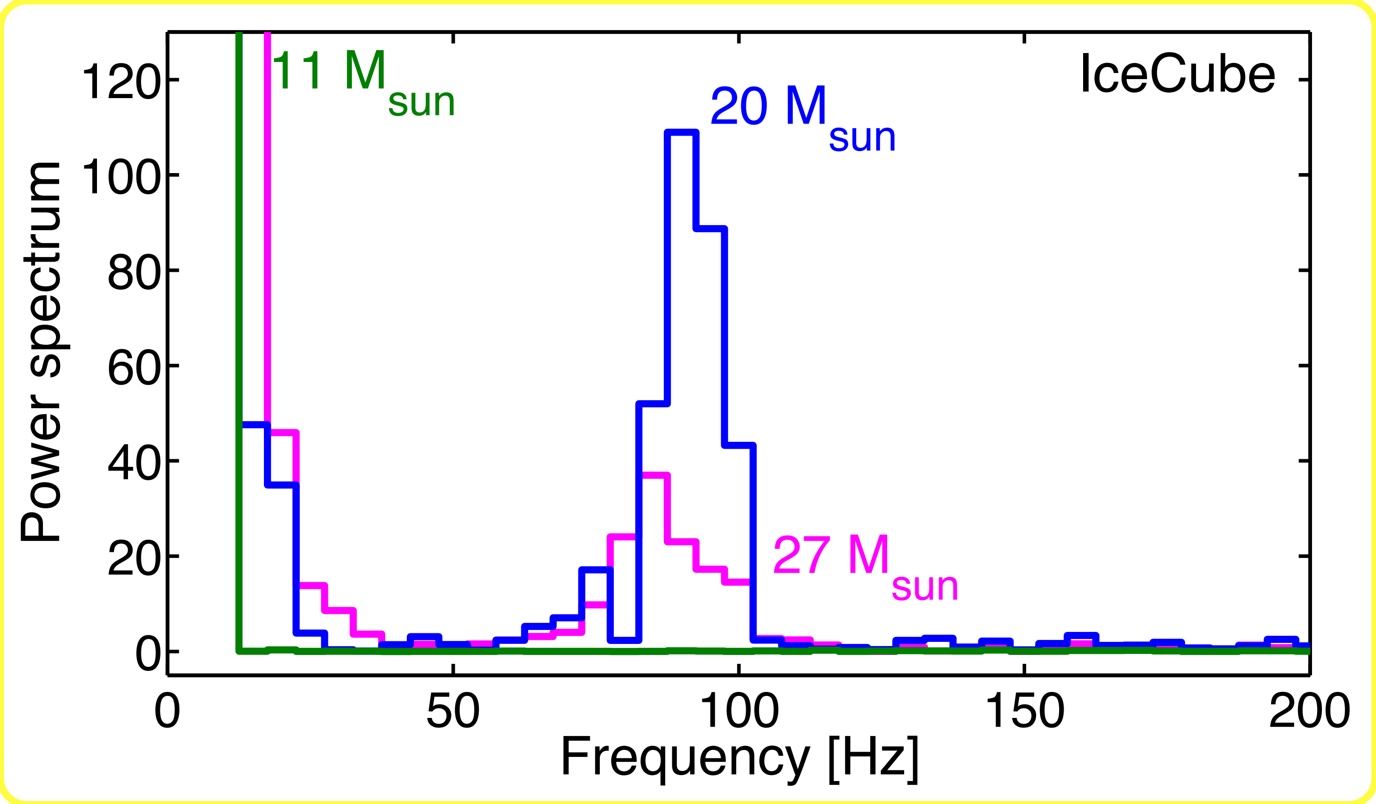
20 M_{sun} SN progenitor:
One SASI episode.

11.2 M_{sun} progenitor:
Large scale convection.

SASI seems to occur for the heavier SN progenitors only.

Power Spectrum of the Event Rate

Power spectrum of the IceCube event rate in [100,300] ms



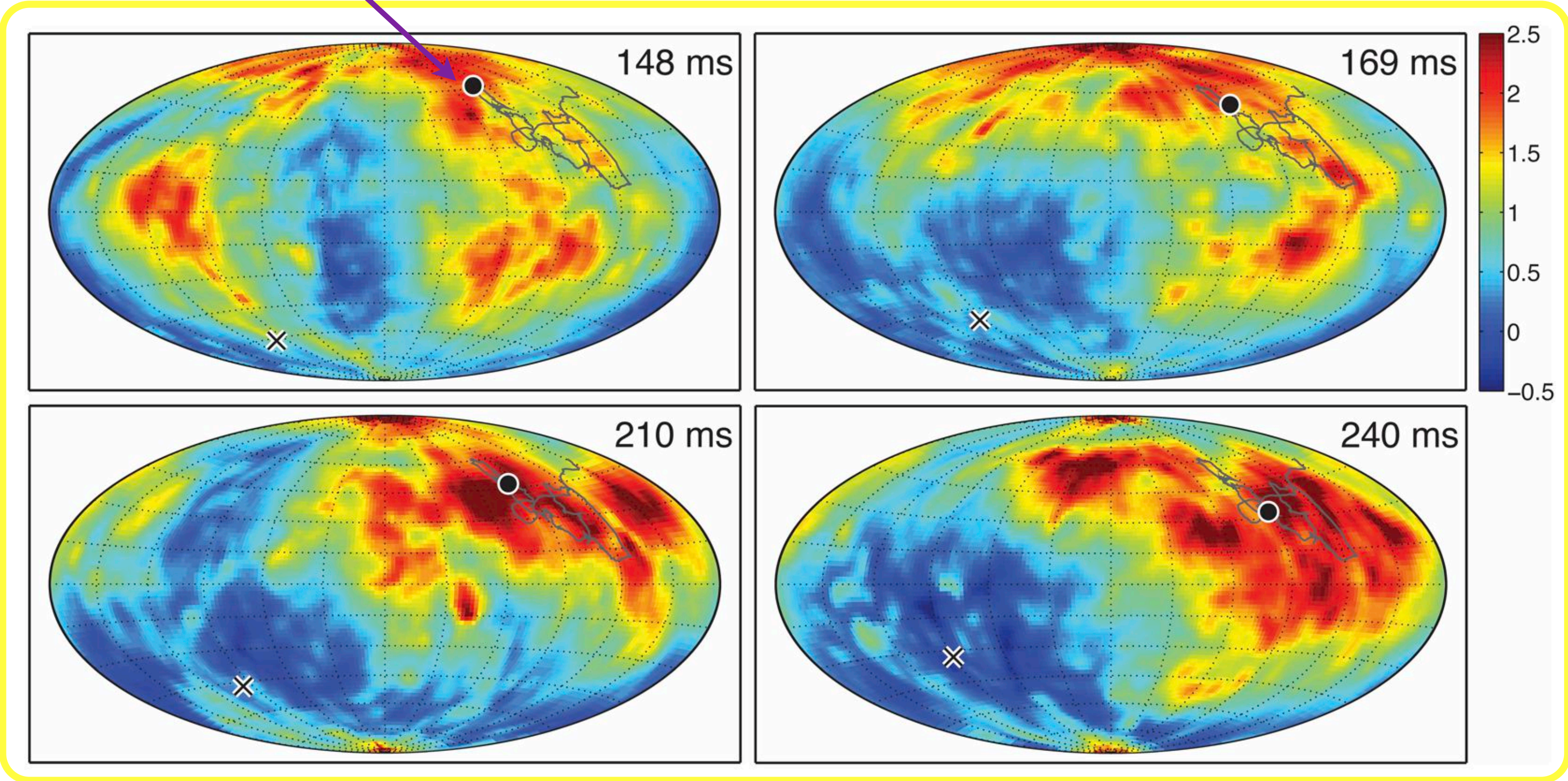
A peak appears at the SASI frequency of ~ 80 Hz for the 20 and 27 M_{sun} SN progenitors.

Lepton-number Emission Self-sustained Asymmetry: A new phenomenon

Lepton-number Flux Evolution

Lepton-number flux for the 11.2 M_{sun} progenitor [$(F_{\nu_e} - F_{\bar{\nu}_e})/\langle F_{\nu_e} - F_{\bar{\nu}_e}\rangle$].

positive dipole direction

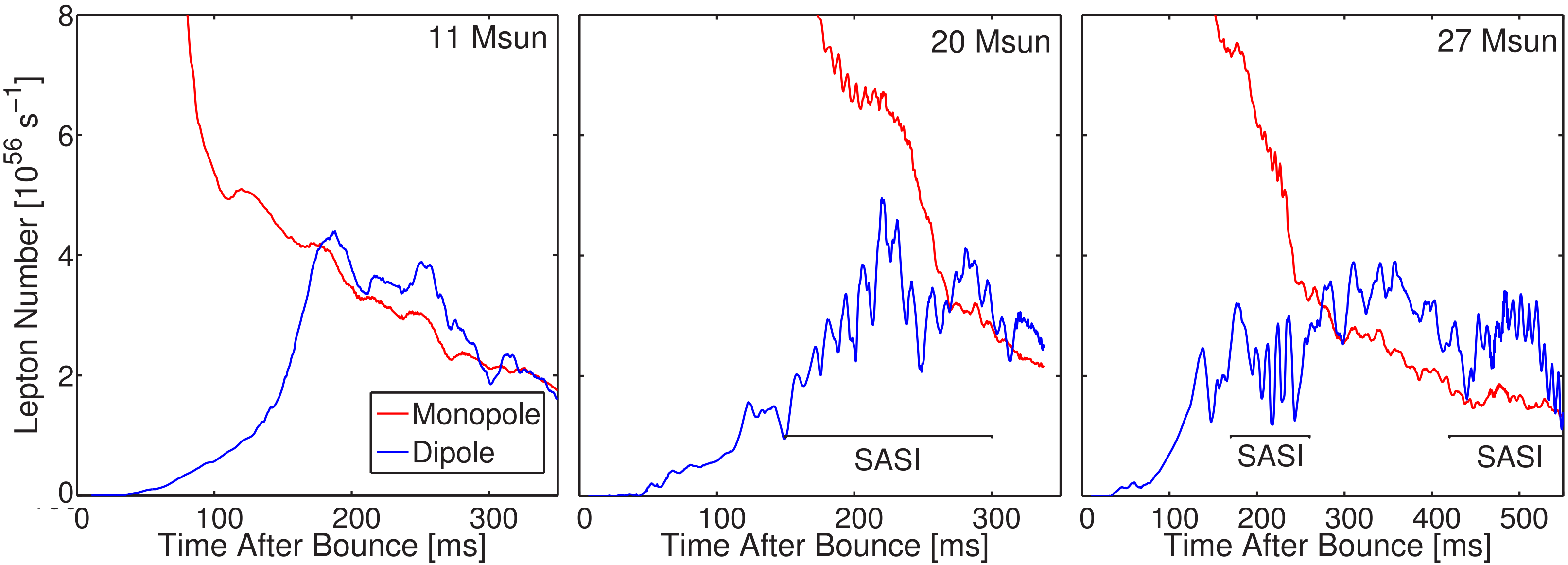


Lepton-number emission asymmetry (**LESA**) is a large-scale feature with **dipole character**.

Once the dipole is developed, its direction remains stable. No-correlation with numerical grid.

Lepton-number Flux Evolution

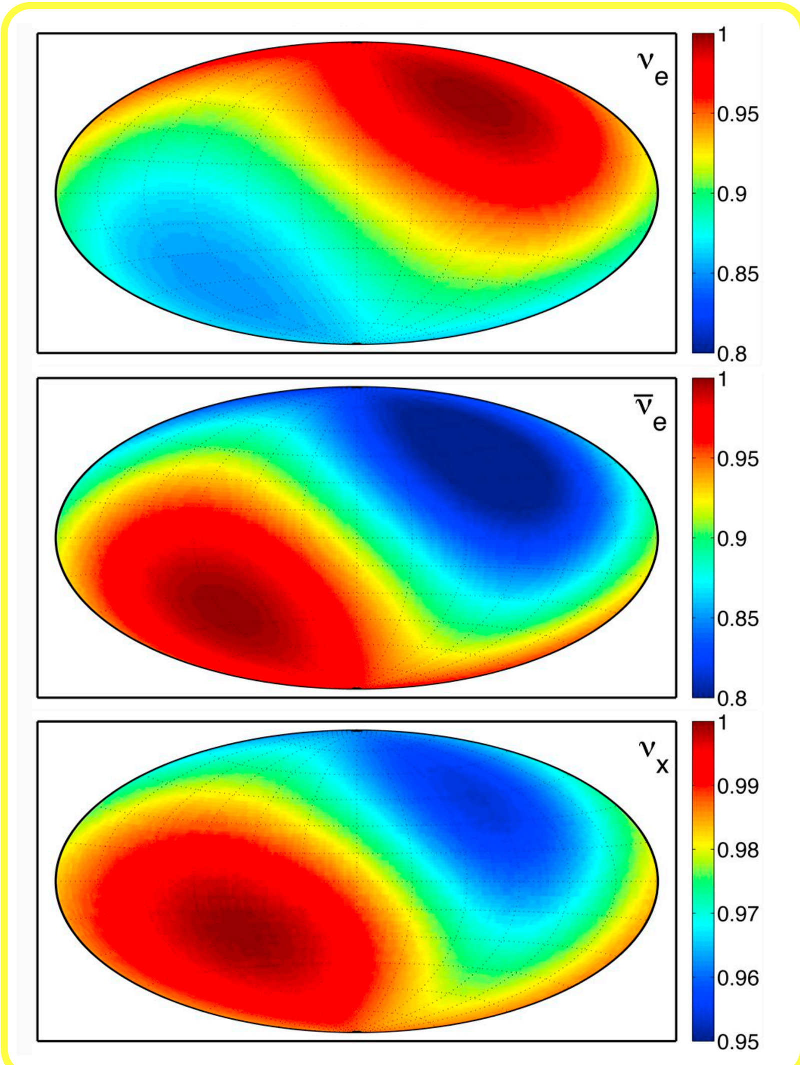
Monopole and dipole of the lepton number flux

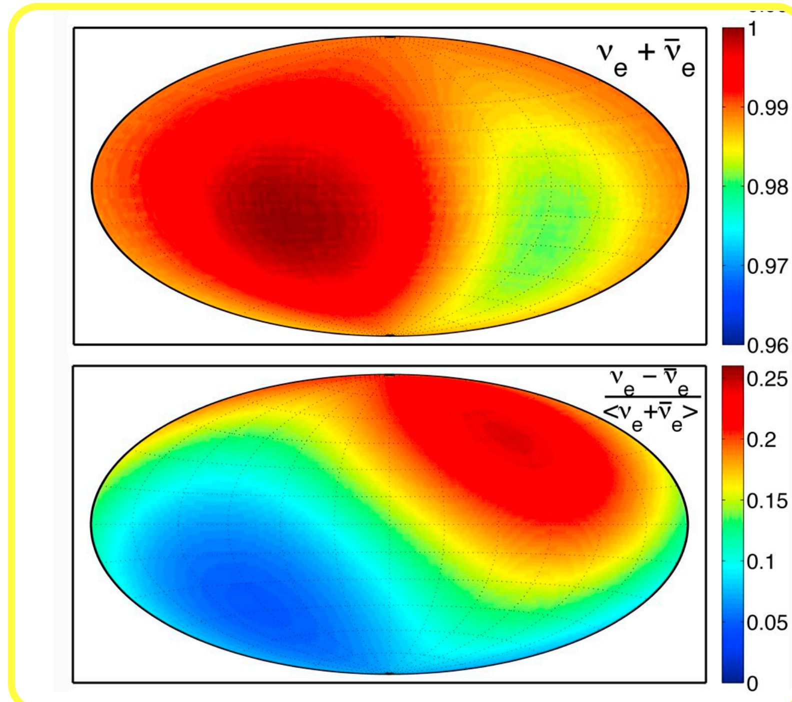


- ★ Monopole evolution strongly depends on the accretion rate and varies between models.
- ★ Maximum dipole amplitude similar in all cases.
- ★ Dipole persists during SASI activity.
- ★ Dipole direction different in each progenitor.

Number Flux Evolution

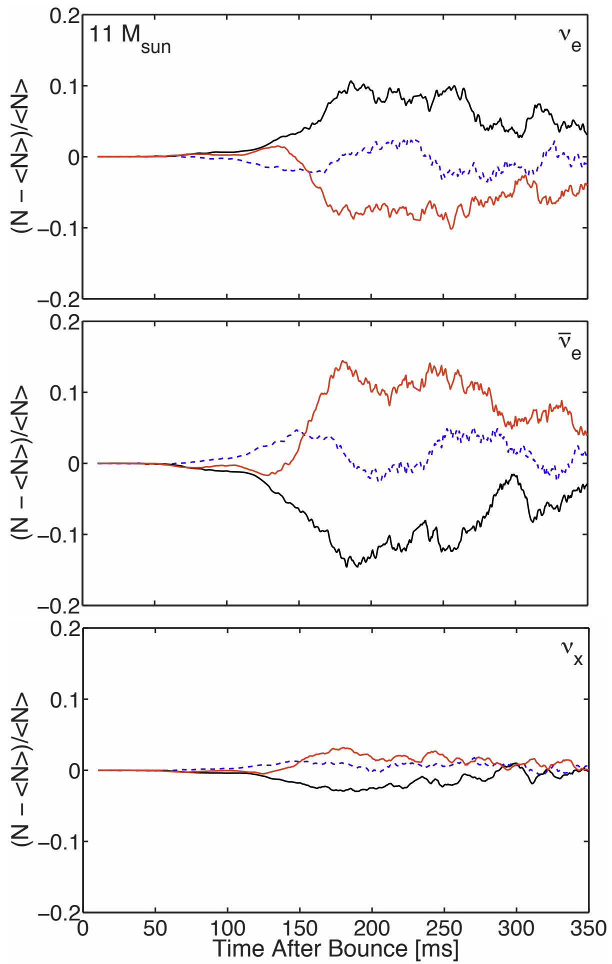
Number flux for the 11.2 M_{sun} progenitor, integrated over [150,250] ms.





Number Flux Evolution

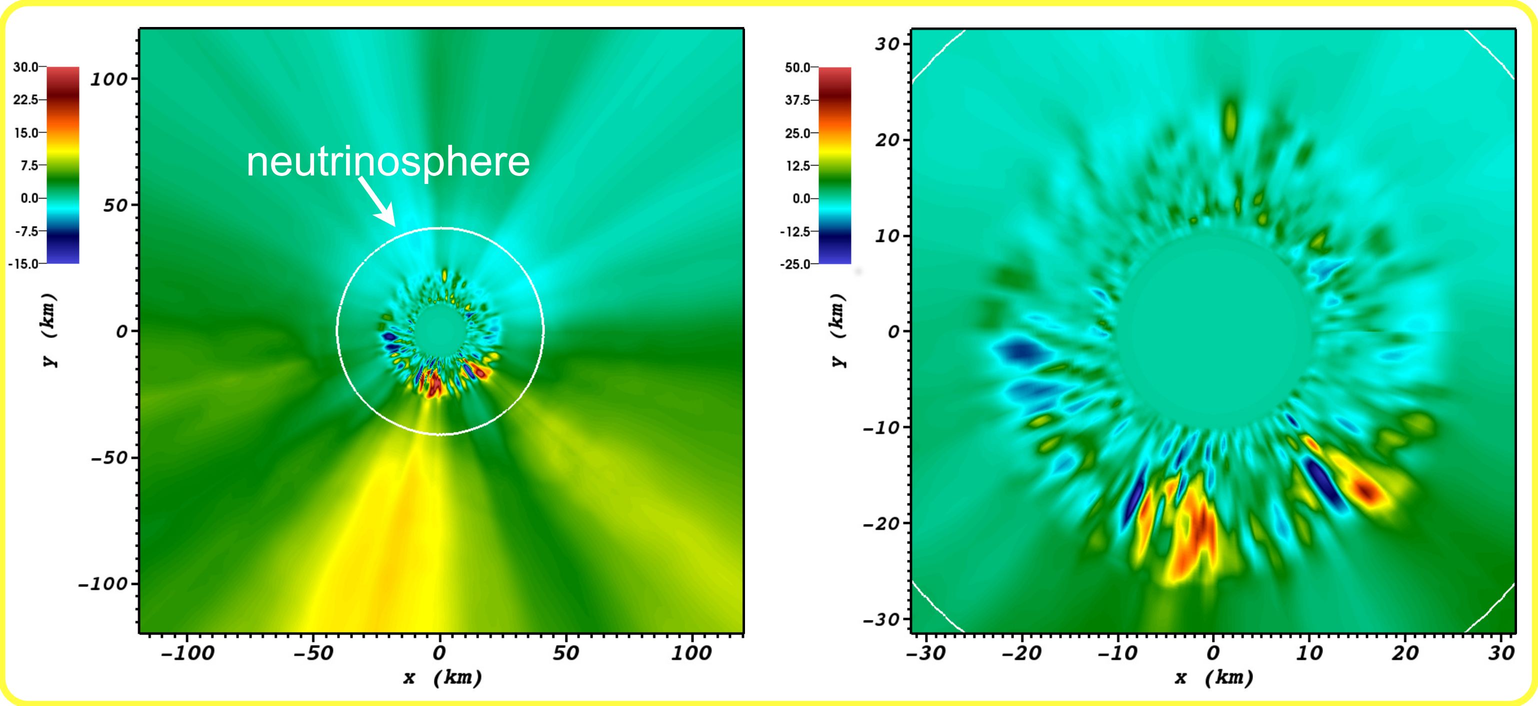
Number flux for 11 M_{sun}



The ν_e and $\bar{\nu}_e$ fluxes exhibit a strong dipolar asymmetry, (anti-)aligned with the lepton-number flux dipole. The ν_x flux is nearly isotropic.

Lepton-number Flux Evolution

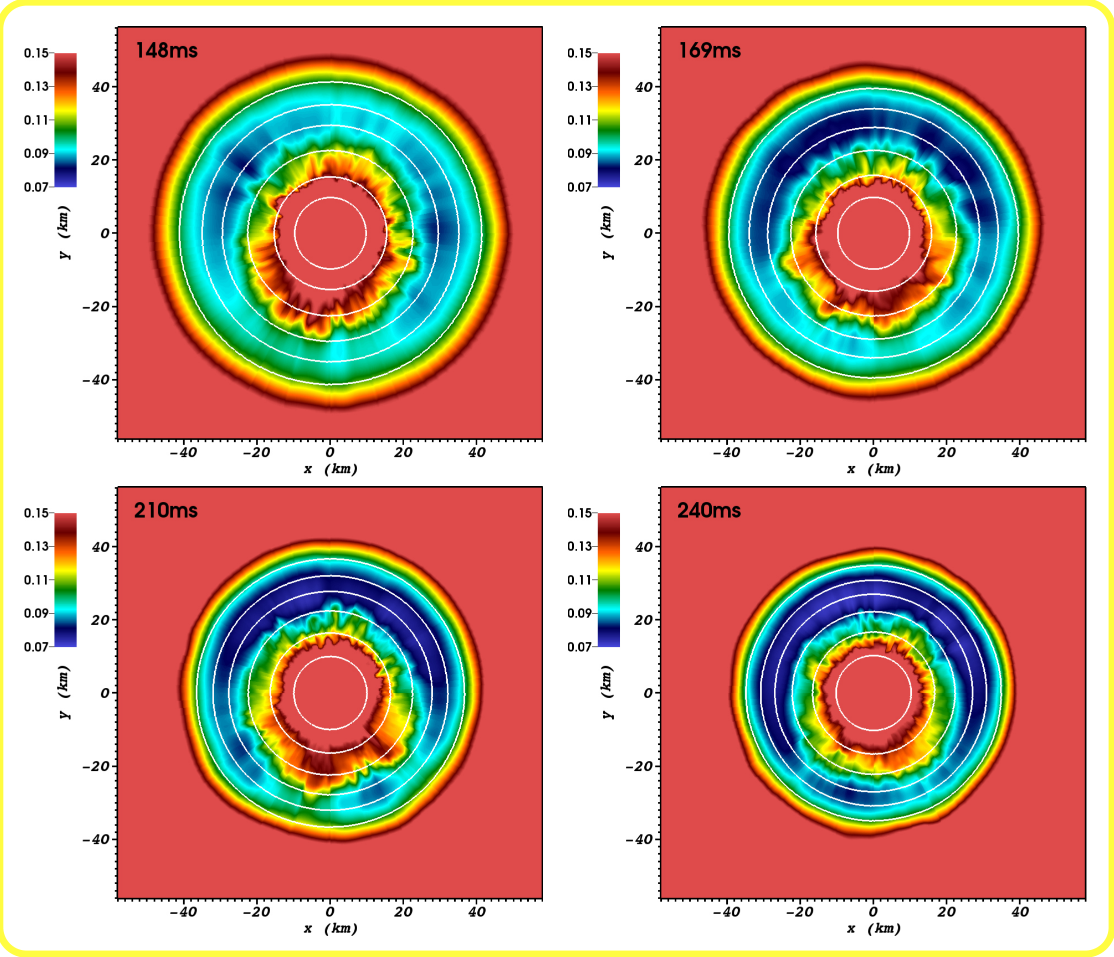
Radial evolution of the lepton-number flux in the the 11.2 M_{sun} progenitor at 210 ms p.b.



Most of the hemispheric difference builds up in the PNS mantle below the neutrinosphere.
PNS convection stronger in the hemisphere of maximal lepton-number flux (bottom direction).

Electron Fraction Evolution

Distribution of the electron fraction in the the 11.2 M_{sun} progenitor.

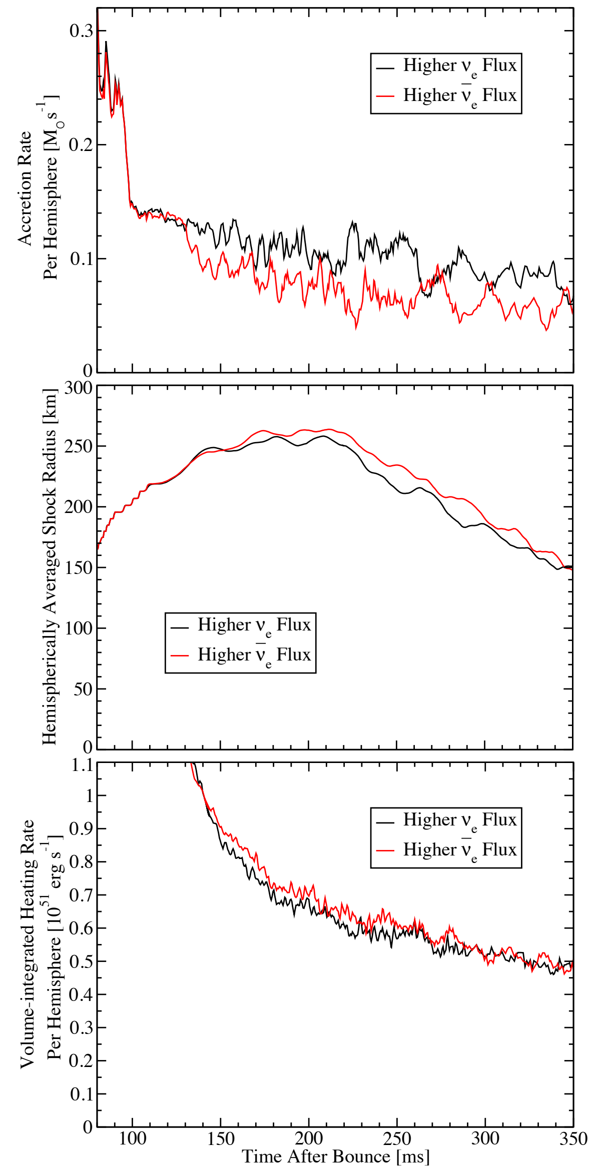


Strongly deleptonized shell in the upper hemisphere (direction of minimal lepton number flux).

Accretion Rate and Shock Radius

Anti-correlation between mass-accretion flow and shock-wave radius.

Neutrino heating is stronger on the side of lower lepton-number flux.



Features of the LESA Phenomenon

★ The initial spherically symmetric state is not stable. LESA grows from any perturbation.

★ LESA is not simply hydrodynamical, but neutrino-hydrodynamical instability in contrast to convection or SASI. First of its kind identified in the SN context.

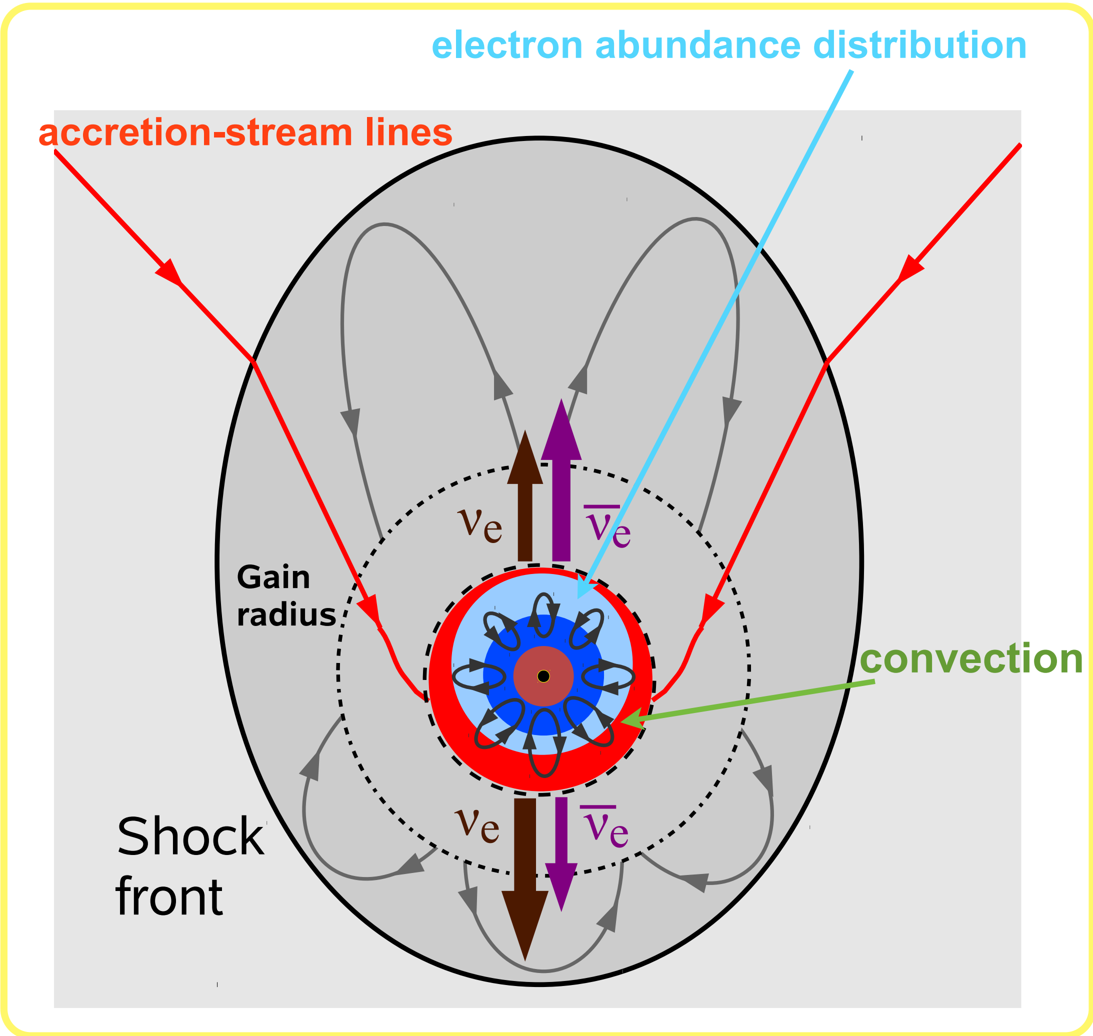
★ LESA mostly builds up below the neutrinosphere.

★ Hemispheric asymmetry of the lepton number flux reaches 20-30% of average values. Sum of neutrino fluxes nearly isotropic.

★ LESA is a self-sustained phenomenon which exists despite convection and SASI activity.

★ LESA is responsible for asymmetric electron fraction distribution, asymmetric accretion rate, asymmetric neutrino heating rate, and dipole deformation of the shock front.

Overall Picture of LESA



Feedback loop consisting of asymmetric accretion rate, asymmetric lepton-number flux, asymmetric neutrino heating rate, and dipole deformation of the shock front.

Implications of the LESA Phenomenon

★ **Nucleosynthesis in the neutrino heated ejecta**: Considerable hemispheric asymmetry of the electron fraction in the neutrino ejecta.

★ **Neutron star kicks**: Asymmetric neutrino emission imparts a recoil on the nascent NS. LESA disfavored as major source for NS kicks.

★ LESA responsible for **angular momentum transfer**, i.e. spin-up of the nascent NS.

★ **Neutrino-flavor conversion:**

- LESA depends on hemispheric asymmetry of neutrino heating rates (modified by oscillations).
- Flavor conversions modify the n/p ratio in the context of nucleosynthesis.
- Directional neutrino-neutrino refraction index.

Conclusions

★ World-wide first 3D SN simulations with detailed neutrino transport available.

★ Neutrinos carry imprints of the SN explosion dynamics.

★ The SN neutrino signal can probe the nature of the hydrodynamical instability.

★ SASI modulations of the neutrino signal detectable in IceCube and Hyper-K.

★ LESA: new neutrino-hydrodynamical instability. Lepton number flux emerges predominantly in one hemisphere.

Thank you

for your attention!