# Constraining $\Lambda\text{CDM}$ and Extensions with the Effective Field Theory of Large-Scale Structures





Based on arXiv:2210.14931 With Pierre Zhang and Vivian Poulin

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### Théo SIMON





# The Effective Field Theory of Large-Scale Structures (EFTofLSS)

To constrain  $\Lambda CDM$  and extensions, there are two main observables: 1. **CMB**  $\rightarrow$  CMB power spectra:  $C_{l}^{TT}$ ,  $C_{l}^{EE}$ ,  $C_{l}^{TE}$ , etc. 2. **LSS**  $\rightarrow$  the galaxy power spectrum:  $P_g(z, k, \mu)$ 

In **linear perturbation theory**, there are two popular ways to use LSS data: 1. Extract information from the full galaxy power spectrum:

 $P_g(z,k,\mu) \simeq$ 

 $b_1$ : bias parameter, f: growth f

2. Redshift Space Distortion (RSE

LSS collaborations conventionally use the second method (+BAO)

$$\left[b_{1}(z) + f\mu^{2}\right]^{2} P_{m}(z,k)$$
[Kaiser '87]

Factor and 
$$\mu = \hat{z} \cdot \hat{k}$$
  
D) information:  $f\sigma_8$ 



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$$C_l^{TT}$$
,  $C_l^{EE}$ ,  $C_l^{TE}$ , etc.

Lack of precision

 $P_g(z,k,\mu) \simeq \left[b_1(z) + f\mu^2\right]^2 P_m(z,k)$ 

[Kaiser '87]

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$$\mu = \hat{z} \cdot \hat{k}$$



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ack of information



### The effective field theory of large-scale structures (EFTofLSS)

The galaxy power spectrum in the framework  

$$P_g(k,\mu) \simeq \left[b_1 + f_{g}(k,\mu) \simeq \left[b_1 + f_{g}(k,\mu) = Z_1(\mu)^2 P_{11}(k) + 2Z_1(\mu) P_{11}(k) \left(d_1 + 2\int \frac{d^3q}{(2\pi)^3} Z_2(q,k-q,\mu)^2 P_{11}(|k-q|) P_{11}(k) + \frac{1}{\bar{n}_g} \left(c_{\epsilon,0} + c_{\epsilon,1}\frac{k^2}{k_{\rm M}^2} + c_{\epsilon,2}f\mu^2\frac{k^2}{k_{\rm M}^2}\right),$$
(Perko et al. (16)

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ork of the EFTofLSS:  $f\mu^2 ]^2 P_m(k) = Z_1(\mu)^2 P_m(k)$ At  $\mathcal{O}(k^2/k_M^2)$ , we go from 1 to 10 free parameters  $\overline{\left(c_{\rm ct}\frac{k^2}{k_{\rm M}^2} + c_{r,1}\mu^2\frac{k^2}{k_{\rm M}^2} + c_{r,2}\mu^4\frac{k^2}{k_{\rm M}^2}\right)}$   $P_{11}(q) + 6Z_1(\mu)P_{11}(k)\int \frac{d^3q}{(2\pi)^3} Z_3(q, -q, k, \mu)P_{11}(q)$ 





# The effective field theory of large-scale structures (EFTofLSS)

The galaxy power spectrum in the framework of the EFTofLSS:  

$$P_{g}(k,\mu) \simeq \left[b_{1} + f\mu^{2}\right]^{2} P_{m}(k) = Z_{1}(\mu)^{2}P_{m}(k)$$

$$At \ \mathcal{O}(k^{2}/k_{M}^{2}), we \ go \ from \ 1 \ to \ 10 \ free \ parameters$$

$$P_{g}(k,\mu) = Z_{1}(\mu)^{2}P_{11}(k) + 2Z_{1}(\mu)P_{11}(k) \left(c_{ct}\frac{k^{2}}{k_{M}^{2}} + c_{r,1}\mu^{2}\frac{k^{2}}{k_{M}^{2}} + c_{r,2}\mu^{4}\frac{k^{2}}{k_{M}^{2}}\right)$$

$$+ 2\int \frac{d^{3}q}{(2\pi)^{3}} Z_{2}(q,k-q,\mu)^{2}P_{11}(|k) - q| P_{11}(q) + 6Z_{1}(\mu)P_{11}(k) \int \frac{d^{3}q}{(2\pi)^{3}} Z_{3}(q,-q,k,\mu)P_{11}(q)$$

$$+ \frac{1}{\bar{n}_{g}} \left(c_{\epsilon,0} + c_{\epsilon,1}\frac{k^{2}}{k_{M}^{2}} + c_{\epsilon,2}f\mu^{2}\frac{k^{2}}{k_{M}^{2}}\right),$$
[Perko et al. '16]  

$$P_{g}(k,\mu) \ can \ be \ determined \ directly \ from \ P_{11}(k) = P_{m}^{lin}(k)$$





# EFTofLSS applied to BOSS data

Multipoles of the galaxy power spectrum, obtained through a **Legendre** polynomials  $(\mathscr{L}_{\ell})$ decomposition:

$$P_g(z,k,\mu) = \sum_{\substack{\ell \text{ even}}} \mathscr{L}_\ell(\mu) P_\ell(z,k)$$

 $\rightarrow$  the two main contributions to  $P_g(z, k, \mu)$  are the **monopole**  $(\ell = 0)$  and the **quadrupole**  $(\ell = 2).$ 



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# EFTofLSS applied to eBOSS QSO data

- $z_{\rm eff} = 1.5$
- 2 skycuts : NGC and SGC



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• 343 708 quasars selected in the redshift range 0.8 < z < 2.2

[eBOSS Collaboration '20, arXiv:2007.08991]







# Determination of the cut-off scale $k_{max}$ of the one-loop prediction

At **one-loop order**, the galaxy power spectrum reads:

$$\begin{split} P_g(k,\mu) &= Z_1(\mu)^2 P_{11}(k) + 2Z_1(\mu) P_{11}(k) \left( c_{\rm ct} \frac{k^2}{k_{\rm M}^2} + c_{r,1} \mu^2 \frac{k^2}{k_{\rm M}^2} + c_{r,2} \mu^4 \frac{k^2}{k_{\rm M}^2} \right) \\ &+ 2 \int \frac{d^3 q}{(2\pi)^3} \, Z_2(q, \mathbf{k} - q, \mu)^2 P_{11}(|\mathbf{k} - q|) P_{11}(q) + 6Z_1(\mu) P_{11}(k) \int \frac{d^3 q}{(2\pi)^3} \, Z_3(q, -q, \mathbf{k}, \mu) P_{11}(q) \\ &+ \frac{1}{\bar{n}_g} \left( c_{\epsilon,0} + c_{\epsilon,1} \frac{k^2}{k_{\rm M}^2} + c_{\epsilon,2} f \mu^2 \frac{k^2}{k_{\rm M}^2} \right), \end{split}$$

One can add the **NNLO terms** (*i.e.*, the dominant two-loop terms):

$$P_{\text{NNLO}}(k,\mu) = \frac{1}{4} b_1 \left( c_{r,4} b_1 + c_{r,6} \mu^2 \right) \mu^4 \frac{k^4}{k_{\text{M}}^4} P_{11}(k)$$

If the contribution of  $P_{\text{NNLO}}(k,\mu)$  becomes **too large,** the one-loop prediction is **not accurate enough**  $\rightarrow$  this determines the **cut-off scale**  $k_{\text{max}}$  of the prediction



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### 2 new EFT parameters





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Additional information

- For **eBOSS**, the error bars of  $\Omega_m$  and  $\sigma_8$  are reduced by a factor  $\sim 2.0$  and  $\sim 1.3$
- For **BOSS**, the error bars of  $\Omega_m$  and h are reduced by a factor ~ 5.4 and ~ 3.2





### LSS data vs Planck



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- eBOSS, BOSS and Planck are consistent at  $\lesssim 1.8\sigma$  on all cosmological parameters
- h is ~  $1\sigma$  lower for eBOSS than for BOSS, while  $\sigma_8$  is ~  $1.5\sigma$  higher
- The h and  $\sigma_8$  Planck values are in-between those of BOSS and eBOSS

 $\rightarrow$  there is no tension between Planck and BOSS/eBOSS





### LSS data vs Planck



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ext-BAO: 6dF & MGS (SDSS) data





# LSS data combined with Planck



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LSS: eBOSS + BOSS + ext-BAO + Pantheon

- Compared to Planck alone, the constraints on  $\Omega_m$  and hare **improved by**  $\sim 30\%$
- $\sigma_8$  and  $A_s$  are not significantly impacted





# **Extensions to** $\Lambda$ **CDM:** curvature density fraction $\Omega_k$

- The EFT analysis significantly improves the constraints on  $\Omega_k$  by ~ 50 % compared to the conventional BAO/ $f\sigma_8$  analysis
- With LSS data only, we find  $\Omega_k$  compatible with zero curvature at  $1.3\sigma$
- The combination of LSS and Planck leads to a **strong constraint** and excludes the (slightly favored) negative values of  $\Omega_k$





### **Extensions to** $\Lambda$ **CDM:** dark energy equation of state $w_0$

- The EFT analysis **improves the constraints** on  $w_0$  by  $\sim 20\%$  compared to the conventional BAO/ $f\sigma_8$  analysis
- With the LSS data only, we find **no evidence for a universe with**  $w_0 \neq -1$
- The addition of LSS data select values of  $w_0$  close to -1, located in the  $2\sigma$  region reconstructed from Planck data





- conventional BAO/ $f\sigma_8$  analysis (  $\sum m_{\nu} = 4.84 eV$ )
- $(\sum m_{\nu} = 0.241 eV)$

### • This analysis **disfavors the inverse hierarchy** at $\sim 2.2\sigma$



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• The EFT analysis **significantly improves the constraints** on  $\sum m_{\nu}$  (by a factor of  $\sim 18$ ) over the

• The LSS constraint derived in this work is only  $\sim 10\%$  weaker than the Planck constraint

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# **Extensions to** $\Lambda$ **CDM:** effective number of relativistic species $N_{\rm eff}$

- parameter
- The value of  $\Delta N_{\rm eff}$  is compatible with the standard model
- The addition of the LSS data **improves** the results of Planck alone



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- The EFTofLSS is a novel method that provides an accurate description of LSS data at a controlled precision
- Constraints from LSS data are competitive with CMB data and their combination improves over Planck alone
- EFTofLSS allows to highlight that there is no tension between current BOSS/eBOSS data and Planck data (but not in tension with weak lensing neither)
- Data are consistent with  $\Lambda$ CDM at  $\leq 1.3\sigma \rightarrow$  Strong constraints on canonical extensions to  $\Lambda$ CDM e.g. LSS+Planck:  $\sum m_{\nu} < 0.093 eV$
- EFTofLSS provides interesting constraints on non-trivial extensions of the  $\Lambda$ CDM model:  $\rightarrow$  see [TS et al. '22, arXiv:2203.07440] for **Decaying Cold Dark Matter**  $\rightarrow$  see [TS et al. '22, arXiv:2208.05930] for Early Dark Energy

### Conclusions





# Thanks for your attention



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### Backup I: EFTofLSS parameters

$$\begin{split} P(k,\mu) &= Z_1(\mu)^2 P_{11}(k) \\ &+ 2\int \frac{d^3q}{(2\pi)^3} Z_2(\mathbf{q},\mathbf{k}-\mathbf{q},\mu)^2 P_{11}(|\mathbf{k}-\mathbf{q}|) P_{11}(q) + 6Z_1(\mu) P_{11}(k) \int \frac{d^3q}{(2\pi)^3} Z_3(\mathbf{q},\mathbf{q},\mathbf{q}) \\ &+ 2Z_1(\mu) P_{11}(k) \left( c_{\rm ct} \frac{\hbar^2}{k_{\rm M}^2} + c_{r,1}\mu \right) \frac{k^2}{k_{\rm M}^2} + c_{r,2}\mu \frac{k^2}{k_{\rm M}^2} \right) + \frac{1}{\bar{n}_g} \left( c_{\epsilon,0} + c_{\epsilon,1} \frac{k^2}{k_{\rm M}^2} + c_{\epsilon,2}f \mu^2 \frac{k^2}{k_{\rm M}^2} \right) \\ \end{split}$$

$$Z_{1}(\mathbf{q}_{1}) = K_{1}(\mathbf{q}_{1}) + f\mu_{1}^{2}G_{1}(\mathbf{q}_{1}) = b_{1} + f\mu_{1}^{2},$$

$$Z_{2}(\mathbf{q}_{1}, \mathbf{q}_{2}, \mu) = K_{2}(\mathbf{q}_{1}, \mathbf{q}_{2}) + f\mu_{12}^{2}G_{2}(\mathbf{q}_{1}, \mathbf{q}_{2}) + \frac{1}{2}f\mu q \left(\frac{\mu_{2}}{q_{2}}G_{1}(\mathbf{q}_{2})Z_{2}(\mathbf{q}_{1}, \mathbf{q}_{2}, \mathbf{q}_{3}) + f\mu_{123}^{2}G_{3}(\mathbf{q}_{1}, \mathbf{q}_{2}, \mathbf{q}_{3})\right)$$

$$+ \frac{1}{3}f\mu q \left(\frac{\mu_{3}}{q_{3}}G_{1}(\mathbf{q}_{3})Z_{2}(\mathbf{q}_{1}, \mathbf{q}_{2}, \mu_{123}) + \frac{\mu_{23}}{q_{23}}G_{2}(\mathbf{q}_{2}, \mathbf{q}_{3})Z_{2}(\mathbf{q}_{1}, \mathbf{q}_{2}, \mu_{123})\right)$$

with

with

$$\begin{split} K_1 &= b_1, \\ K_2(\mathbf{q}_1, \mathbf{q}_2) &= b_1 \frac{\mathbf{q}_1 \cdot \mathbf{q}_2}{q_1^2} + b_2 \left( F_2(\mathbf{q}_1, \mathbf{q}_2) - \frac{\mathbf{q}_1 \cdot \mathbf{q}_2}{q_1^2} \right) + b_4 + \text{perm.}, \\ K_3(k, q) &= \frac{b_1}{504k^2q^3} \left( -38k^5q + 48k^3q^3 - 18kq^5 + 9(k^2 - q^2)^3 \log \left| + \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \right| \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(3k^4 - 14k^2q^2 + 3q^4) + 3(k^2 - q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(k^2 + q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(k^2 + q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(k^2 + q^2) \right) \\ &= \frac{b_3}{758k^3q^5} \left( 2kq(k^2 + q^2)(k^2 + q^2) \right) \\ \\ &$$

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### **10 parameters**

**4 parameters**  $b_i$  (i = 1, 2, 3, 4) to describe the galaxy bias which arises from the one-loop contributions.

3 parameters corresponding to **counterterms** ( $c_{ct}$  linear combination of a higher derivative bias and the dark matter sound speed, while  $c_{r,1}$  and  $c_{r,2}$  are the redshift-space counterterms).

3 parameters which describe stochastic

terms.



### Backup II: test against simulations



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- EZmock: mocks that are built to simulate eBOSS observational characteristics
- We find that up to  $k_{\rm max} = 0.24 h \, {\rm Mpc^{-1}}$ , the best-fit values of the cosmological parameters are shifted with respect to the truth of the simulations by  $\lesssim 1/3 \cdot \sigma$



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# Backup III: variation of $n_s$ and $\omega_b$



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- We impose a uninformative large flat prior on  $n_s$ , while we impose a BBN Gaussian prior on  $\omega_b$
- The variation of  $\omega_b$  within the BBN prior has a negligible impact on the cosmological results: we have a relative shift of  $\lesssim 0.04\sigma$
- The variation of  $n_s$  within a uninformative large flat prior leads to a relative shift  $~\lesssim 0.4\sigma$

![](_page_23_Picture_8.jpeg)

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# **Backup IV:** dark energy equation of state $w_0 \ge -1$

- negligible way, while it remains globally stable for the LSS + Planck
- For these analyses,  $\Delta \chi^2 = 0$  with respect to  $\Lambda$ CDM, since we obtain best-fit values of  $w_0 = -1$

![](_page_24_Figure_3.jpeg)

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• One can see that this new prior shifts the 2D posteriors inferred from the LSS data in a non-

![](_page_24_Picture_9.jpeg)